

KEY TO METEOROLOGICAL RECORDS DOCUMENTATION NO. 5.313

Catalog of Meteorological Satellite Data-ESSA 3 Television Cloud Photography

Part 2



PURPOSE

The Key To Meteorological Records Documentation Series has been established to provide guidance information to research personnel making use of climatological data.

Frequently users of such data have found it necessary to spend a great deal of time establishing whether the criteria for observing or computing various elements have changed over the period of record or in what form the data are available.

It is therefore hoped that the presentation of this series may not only conserve valuable time but may have a direct influence in improving the accuracy of research results.

EARLIER TIROS AND ESSA DATA CATALOGS

Earlier catalogs of TIROS and ESSA Meteorological Satellite data are available in this series, as follows:

- No. 5.31 "Catalogue of Meteorological Satellite Data TIROS I Television Cloud Photography", published in 1961, price 70 cents.
- No. 5.32 "Catalogue of Meteorological Satellite Data TIROS II Television Cloud Photography", published in 1963, price 20 cents.
- No. 5.33 "Catalogue of Meteorological Satellite Data TIROS III Television Cloud Photography", published in 1962, price 70 cents.
- No. 5.34 "Catalogue of Meteorological Satellite Data TIROS IV Television Cloud Photography", published in 1963, price \$1.00.
- No. 5.35 "Catalogue of Meteorological Satellite Data TIROS V Television Cloud Photography", published in 1964, price \$1.75.
- No. 5.36 "Catalogue of Meteorological Satellite Data TIROS VI Television Cloud Photography", published in 1964, price \$2.00.
- No. 5.37 "Catalogue of Meteorological Satellite Data TIROS VII Television Cloud Photography Part 1 June 19, 1963, to December 31, 1963", published in 1965, price \$1.25.
- No. 5.37 "Catalogue of Meteorological Satellite Data TIROS VII Television Cloud Photography Part 2 January 1, 1964, to June 30, 1964", published in 1965, price \$1.00.
- No. 5.37 "Catalogue of Meteorological Satellite Data TIROS VII Television Cloud Photography Part 3 July 1, 1964, to December 30, 1964", published in 1965, price \$1.00.
- No. 5.37 "Catalogue of Meteorological Satellite Data TIROS VII Television Cloud Photography Part 4 January 1, 1965, to December 31, 1965", published in 1966, price 45 cents.
- No. 5.38 "Catalogue of Meteorological Satellite Data TIROS VIII Television Cloud Photography Part 1 December 21, 1963, to June 30, 1964", published in 1965, price \$1.00.
- No. 5.38 "Catalogue of Meteorological Satellite Data TIROS VIII Television Cloud Photography Part 2 July 1, 1964, to December 31, 1964", published in 1965, price \$1.00.
- No. 5.38 "Catalogue of Meteorological Satellite Data TIROS VIII Television Cloud Photography Part 3 January 1, 1965, to August 31, 1965", published in 1966, price 45 cents.
- No. 5.39 "Catalogue of Meteorological Satellite Data TIROS IX Television Cloud Photography Part 1 January 22, 1965, to April 30, 1965", published in 1966, price \$1.75.
- No. 5.39 "Catalogue of Meteorological Satellite Data TIROS IX Television Cloud Photography Part 2 May 1, 1965, to July 26, 1965", published in 1967, price \$1.25.
- No. 5.310 "Catalogue of Meteorological Satellite Data TIROS X Television Cloud Photography Part 1 July 2, 1965, to September 30, 1965, in press.
- No. 5.311 "Catalogue of Meteorological Satellite Data ESSA 1 Televis on Cloud Photography Part 1 February 3, 1966, to March 31, 1966", published in 19.6, price \$1.00.



U.S. DEPARTMENT OF COMMERCE Alexander B. Trowbridge, Secretary ENVIRONMENTAL SCIENCE SERVICES ADMINISTRATION Robert M. White, Administrator

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KEY TO METEOROLOGICAL RECORDS DOCUMENTATION NO. 5.313

Catalog of Meteorological Satellite Data-ESSA 3 Television Cloud Photography

Part 2

January 1, - March 31, 1967

PREFACE

This issue is number two in a series of catalogs describing the television cloud photographs obtained by ESSA 3 meteorological satellite. The material in this issue covers the period of January, February and March of 1967. The preparation of this catalog was expedited by computer usage, thus it is being published prior to part one in the series which is being prepared manually. The maps and listings generally follow the pattern established in the "Catalogue of Meteorological Satellite Data - TIROS I Television Cloud Photography," published as No. 5.31 in the Weather Bureau Series Key to Meteorological Records Documentation. However, digitized cloud maps replace the manually constructed nephanalyses which have been used in previous catalogs. There are two maps for each day; the first polar stereographic projection in each set is for the Northern Hemisphere and the second for the Southern Hemisphere. These maps are arranged according to pass day.

Documentation Section
National Environmental Satellite Center

USING THE CATALOG

- 1. Scan the digitized cloud maps for the dates and areas of particular interest.
- 2. When a picture of interest is located, use the appropriate track map, figures 2 and 3, to determine the track numbers by which to order.
- 3. When the track number has been determined, along with the date on which the pass was acquired, the film can be ordered from the National Weather Records Center, Asheville, N. C., at a cost of \$6.50 per 100 foot reel.
- 4. Information for a particular pass, such as start time, subsatellite point (the point where the local vertical through the satellite intersects the earth's surface), and track number, may be determined by using the tabulated listings preceding the digitized cloud maps.

CATALOG OF METEOROLOGICAL SATELLITE DATA

ESSA 3 TELEVISION CLOUD PHOTOGRAPHY

PART 2 - January 1, 1967 to March 31, 1967

The ESSA 3 meteorological satellite, launched on October 2, 1966 by the National Aeronautics and Space Administration, is the third Environmental Survey Satellite in the TIROS Operational (TOS) System. ESSA 3 was put into a nearly circular, sun-synchronous polar orbit, approximately 892 statute miles above the earth's surface with an apogee of 925 statute miles and a perigee of 859 statute miles. The orbit is inclined 79° (Retrograde) to the equatorial plane with an orbital period of 114.5 minutes which corresponds to slightly more than 13 passes around the earth each day.

ESSA 3 is a TIROS type satellite in a "cartwheel" configuration which allows earthoriented picture coverage. The spacecraft is spin-stabilized and magnetically torqued to a
wheel attitude, so that the spin axis is normal to the plane of the orbit and the radially
mounted cameras view the earth once each spacecraft revolution.

The cameras in the ESSA 3 satellite are one inch diameter vidicons (television camera tube) of the Advanced Vidicon Camera System (AVCS) type. Each picture covers an area of 4,000,000 square miles. One camera provides full global coverage, therefore, two cameras provide full system redundancy.

Each of the cameras can independently take pictures and store them for later play back to the Command-and-Data-Acquisition (CDA) stations located at Wallops Station near Chinco-teague, Va., and Gilmore Creek, Fairbanks, Alaska

Pictures, taken every 260 seconds, are obtained in sequences of twelve frames to a pass. Each frame is electronically gridded by computer with latitude and longitude lines and geographic outlines merged with the picture. An identification legend underneath each picture provides the following information: year, month, day, hour, minute and second of picture taking time; track and zone numbers; station initial; satellite number; mode; camera; latitude of subsatellite point; latitude spacing of grids in degrees; longitude of subsatellite point; longitude spacing of grids in degrees; orbital pass number; frame number; and sun glint position.

In the example of figure 1, the legend indicates the picture was taken on orbital pass number 1675 by ESSA 3 and that it is a TAPE picture taken by camera 2 on February 12, 1967 at 2038072. The frame is number four in the sequence and was acquired at Fairbanks, Alaska. The picture is located in the area of track 4 (TABLE 1) zone 58 (not referenced), with the subsatellite point of the picture located at 25°N and 105°W. The grids are spaced 5° both in latitude and longitude and sun glint is shown in the picture at 19°N and 109°W.

For ESSA 3 the system of "TRACKS" is used to identify the general geographic areas over which picture sequences were taken. These tracks are shown on the locator maps (figs. and listed in the track table (TABLE 1) which appear in this catalog immediately following the tabulated listings. The track boundaries are parallel to the path followed by the subsatellite point during the daylight portion of each orbital pass. The width of each track is exactly equal to the amount the earth rotates between consecutive passes. Thus, in general, the subsatellite point makes one traverse along each track each day (from midnight to midnight, GMT). The track boundaries are fixed permanently for the life of ESSA 3, and are arbitrarily referenced to the point 0° latitude, 0° longitude, with track numbers increasing westward.

Since the ESSA 3 orbital period of 114.5 minutes is not an exact submultiple of 24 hours, the satellite starts 12 and 13 passes on alternate days. This results in track 13 being much narrower than the others and, thus, it is not traversed everyday. There are also days on which one particular track is traversed just after GMT midnight, and again later the same day just before GMT midnight.

On the film reels, picture sequences are grouped together by track number and assembled

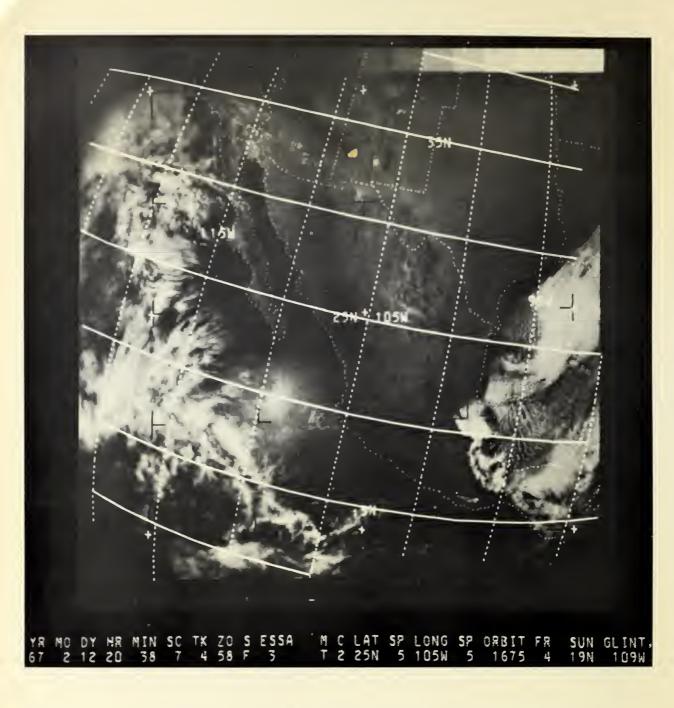


Figure 1 - Example of ESSA 3 Cloud Photography

in chronological order. Track numbers are stamped on the film immediately preceding the sequences to which they apply.

The sixteen pages of tabulated listings give descriptive information about the pictures obtained by ESSA 3. In the listings each picture sequence is described by one line, and the sequences are grouped together by date and track number. This is the order in which they appear on the film reel. The column headings and entries have the following meanings:

PASS: Orbital pass number on which the pictures were taken. On the digitized cloud maps in this catalog, these numbers appear beneath the date in the lower right corner of the page.

TRACK: Defines the geographic area which contains the path of the subsatellite point. The entries refer to the numbered areas outlined on figures 2 and 3 in this catalog.

C: Camera number indicated in the legend.

1 = Camera number one

2 = Camera number two

S: Command-and-Data-Acquisition (CDA) station which received the sequence.

W = Wallops Station, Va.

F = Fairbanks, Alaska

ALONG: Longitude of ascending node (equator crossing)

MODA: Month and day on which the pictures were taken.

HRMNSE: Time in hours, minutes and seconds of the first picture

taken.

SUBSATELLITE POINT PICTURE LOCATION: Indicates the latitude in degrees of the subsatellite point for each of the twelve frames in the orbital pass sequence. First latitude column corresponds to picture frame No. 12.

Beginning with this catalog, a series of digitized computer mapped photographs will replace the manually constructed nephanalyses previously used in catalogs of this series. These DIGITIZED CLOUD MAPS are prepared by means of a high speed digital computer program. In this process the signals comprising the picture taken by the satellite are assigned numerical values to indicate the relative brightness of each element. These data are brightness normalized, earth-located and repositioned on standard map projections. Magnetic tapes are produced for input to a cathode ray film display device.* The photographic product resulting from this computer processing consists of 16 film images; eight polar stereographic octants, 6 Mercator segments and two reduced resolution polar hemispheric chips for each day. The digitized cloud maps used in this catalog are the reduced resolution nothern and southern polar hemispheric chips in which the latitude and longitude lines are spaced at 10°. In some instances, blank areas and mislocated clouds appear on the digitized cloud maps which are irregularities of the computer operation and should be disregarded.

^{*} A detailed description of the procedures used to produce these digitized cloud mosaics follows this section.

The track number and date should be utilized as the primary identification when ordering film. The track number for each sequence is shown on the equatorial line of the digitized cloud maps, and the date appears in the lower right corner of the page along with the inclusive pass numbers for the day.

The ESSA 3 master films are deposited at the National Weather Records Center (NWRC), Environmental Science Services Administration, Federal Building, Asheville, North Carolina 28801. Persons or institutions desiring copies may order them from the NWRC in the form of 35 mm. positive transparencies for projection or 35 mm. duplication negatives from which opaque prints can be made. Two days of pictures, including the digitized computer maps for each day (16 film chips per day), are stored on 100 foot reels, and should be ordered by track number and date. The digitized computer maps are also available separately on 100 foot reels, with each reel containing approximately 22 days of the film chips. These should also be ordered by track number and date. Orders must be placed for one or more complete reels, at a cost of \$6.50 each, as it is not now possible to furnish copies of individual frames or to provide enlargements or other picture formats. The track numbers of the other tracks contained on the 100 foot reel can be obtained from the NWRC prior to ordering the film. All copies will be furnished with sprocket holes since the necessary film emulsion is available only in this form.

Available also from the NWRC, for the cost of reproduction, are reels of microfilmed nephanalysis charts for TIROS I thru TIROS IX.

The digitized mosaics of cloud pictures in this catalog, which replace the nephanalyses of the earlier catalogs, should facilitate retrieval of specific photographs for detailed study and for research purposes. The process by which these mosaics were produced is described herein.

1. INTRODUCTION

Nephanalyses prepared manually from typical orbital picture mosaics are expressable in terms of 8- to 10-min. teletypewriter messages--the equivalent of only 30,000 to 50,000 binary bits per mosaic. However the mosaic itself, expressed in digital form at 2-mile resolution with 4-bit brightness elements, amounts to more than 6 X 10⁶ bits. Thus the data bulk is reduced about two orders of magnitude with hand methods. Although this is probably not a good indicator of information loss, developmental studies over the past several years do suggest that the images contain a far greater quantity of information than was previously extracted for operational use.

The purpose of this description is to acquaint the current and potential users of satellite cloud picture data with the computer products which are now operational. The procedure is described for producing full resolution rectified orbital mosaics for direct visual uses.

2. VIDEO DATA DIGITIZING AND FORMATTING

The first step in the process involves the digitizing of incoming video data which are received via microwave link from the Command and Data Acquisition (CDA) Stations. Much of the specialized equipment originally employed in digitizing Nimbus video data has been used in the present system. However, the presently operational Digital Format System (DFS) has undergone substantial logic changes so that it now operates in some seven different modes.

Briefly, in the case of video data from ESSA satellites, the analog (FM) signals are passed through a discrimination stage so that each raster line appears as a time variable d.c. voltage train. Each such voltage train is sampled repeatedly and converted to a corresponding series of 6-bit digits (bytes). The sample population is adjusted to 800, so that the 800 earth-viewing raster lines produce a 640,000-element array.

The sample bytes are packed into 24-bit words and transferred in parallel directly into the memory of a medium scale computer with the aid of certain interface control signals. Additional "interrupt" signals are provided to the computer through synchronizing detectors to arrange the incoming data into "frames" with reference to the beginning of each video raster and with reference to each horizontal raster line within the image. The edited raw digital data are finally produced on magnetic tape ready for further computer processing as described below.

3. DIGITAL RECTIFICATION PROCEDURES

The program for video data ingestion on the Control Data Corporation CDC-924 computer includes only minimum monitoring and conditioning of the raw sample values. Marker words are added in response to pulses received from the synchronizing detectors and other key format words are also added during the time interval between the acceptance of the digitized video data and the copying of the raw information onto digital video tape. There is need for further conditioning as the data are made available to the large-scale CDC-6600 computer.

^{*} This description is a condensed and edited version of "Operational Processing of Satellite Cloud Pictures by Computer," by C.L. Bristor, W.M. Callicott, and R.E. Bradford, in the MONTHLY WEATHER REVIEW, 94(8), August 1966, pp 515-527.

First, certain samples at the picture edges are removed by a preliminary edge-cropping procedure. The amount of data cropping in a raster row is indicated in an identification word which is inserted as the first word of each retained video raster row. (Plans are projected for the removal of noise bursts and fiducial marks and the interpolative insertion of realistic substitute video information.)

The data are further conditioned by a brightness calibration process. In order to make the individual frames uniform in brightness response, each 16 x 16 raster sample subset is calibrated for gray scale value. The raw video data range in gray scale from 0 to 63. A "zone argument" is determined for each subset and, from the proper zone calibration table, a replacement gray scale value is selected as a function of the original sample brightness. Thus "normalized" brightness fields are provided for the rectification and analysis programs. The final normalized gray scale values range from 0 to 14.

4. RECTIFICATION

The video data are now ready for transformation to arrays oriented to a map projection. The actual rectification process is, in effect, the assignment of raster elements to earth coordinates for replotting. Latitude and longitude values computed for every 32d raster point along every 32d raster row provide an open lattice of earth locations which are used as benchmarks for interpolating the locations of individual raster samples. This lattice was selected because: (a) the computation required for establishing a latitude-longitude value for every point is excessive, (b) the interval between benchmark points approximates a straight line on earth thus minimizing linear interpolation error, and (c) the value 32 is a binary factor permitting computer interpolation shortcuts. Finally, the open lattice latitude-longitude values are converted to map I and J coordinates in a square mesh map overlay. A benchmark table of map I and J values is generated for each video frame processed (reflecting its unique perspective).

5. ORBITAL MOSAICS

The video data are rectified to polar stereographic and Mercator map projections with grid mesh size commensurate with the original resolution. In addition, the polar stereographic array is a binary superset of the commonly used numerical weather prediction (NWP) grid system. This array is divided, step-wise, into two sets (see fig. 2). One is an 8 X 8 set of the NWP system and the other is an 8 X 8 set of the first thus providing a 64 X 64 superset. A linear map resolution of approximately 2 n.mi. at the equator and 4 n.mi. at the pole is realized.

The rectification or replotting logic requires memory space for all video data samples within the map array region being viewed. Video data are stretched or distorted over the array as a function of the map projection mathematics. This stretching is more severe in image areas where foreshortening is evident. Since the complete 64 X 64 polar array has 4096 X 4096 grid squares per hemisphere, methods for array storage compaction are required. An orbital swath of data requires a minimum of 2,464,000 grid squares, which exceeds the available high speed computer memory capacity. The polar array grid square area required for a swath is, therefore, broken into parts called subregion arrays. Each subregion is an irregular skew-shaped array containing those grid square "bins" in which imagery is stored. Because the bins are arranged serially in the computer, a two-dimensional indexing scheme is required for random access to the data (see fig. 3). Even so, the irregular arrays require memory space for 224,000 samples so only computers with large-scale memory capacity can be used efficiently for the rectification problem.

The Mercator grid square array contains 32 mesh intervals per degree of longitude. This fixed dimension array consists of 1600 columns and 2000 rows. The extreme right hand column is aligned with a meridian selectable in even 5° intervals and the top grid row overlies the top (also selectable) latitude limit (see fig. 4). The indexing problem is not as severe as that for the polar array because the subregions are rectangular.

The extreme volume of grid-square data involved in full resolution rectification stresses the need for efficient programing and for the elimination of redundant processing. In the

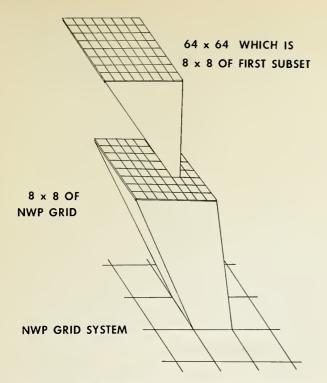


Figure 2 - The method is illustrated whereby the NWP mesh is subdivided to obtain a 64 x 64 superset. The two-step breakdown facilitates computer handling.

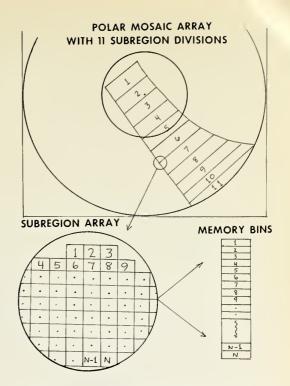


Figure 3 - The need for indexing logic is indicated by the sequential storage of skewed data trains.

case of Advanced Vidicon Camera System coverage of the ESSA satellites, there is generally considerable frame-to-frame overlap and significant overlap between orbital passes - particularly at middle and high latitudes. So the overlap problem is alleviated through the use of internal cropping logic.

The satellite's orbital characteristics and the camera shutter time schedule establish a set of cropping tables which are adjustable to process any desired picture portions. Each frame is cropped uniquely. This permits cropping to minimize sun glint problems and to avoid using imagery having poor brightness response or extreme foreshortening. Central raster lines provide better source imagery from the standpoint of brightness characteristics and resolution, but near the equator where overlap between passes is minimal, the center raster lines contain the sun glint. Where possible (at higher latitudes) sun glint is removed by the use of overlapping imagery (see figs. 5 and 6). Parameters are supplied to the rectification program in a format permitting asymmetric cropping, which eliminates processing of data outside the designated area. Minimum redundant processing is allowed to guarantee continuity within an orbital swath and to avoid gaps between passes.

Care is taken to eliminate data voids within the mosaic images thereby providing greater eye appeal and guaranteeing continuity for follow-on processing. There are areas in the rectified mosaics where the map resolution exceeds that of the data source. This is particularly true near the equator on the polar array and at high latitudes on the Mercator array. Since the coarser resolution video samples apply almost as well to adjacent higher resolution map squares, a simple areal average is used to fill voids in the polar array. For convenience the Mercator array is filled by a bi-directional row scan which propagates adjacent values into voids. Both processes create smooth mapped images where voids occur, preserving a continuous picture image. Examples of full resolution orbital mosaics are shown in figures 7 and 8.

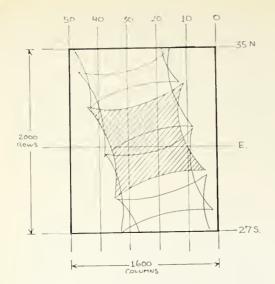


Figure 4 - The Mercator mosaic map coverage is indicated (heavy border) along with the data array dimensions.

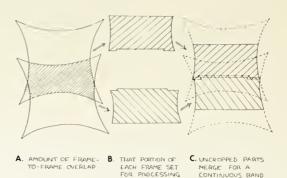


Figure 5 - The frame-to-frame overlap permits arbitrary cropping.

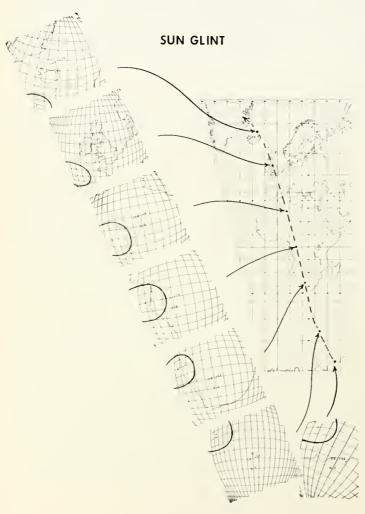


Figure 6 - Sample earth locator grids are shown as insets on a typical mapped pass for a sun synchronous orbit. It is to be noted that sun glint occurs in all frames.

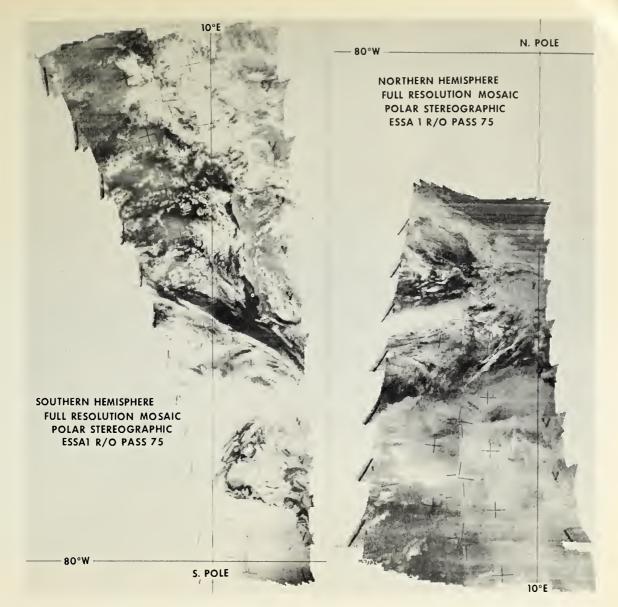


Figure 7 - Approximately 40 ESSA 1 pictures are combined into full resolution polar stereographic mosaics. Interpolative filler has been applied where foreshortened imagery has been "stretched" to the map array. Interim brightness calibration utilizing tables from figure 12 has not yet been applied, so "seams" are still apparent in certain overlap areas.

6. MONTAGE GENERATION

As the mosaics for adjacent orbital swaths are generated, they are blended into a composite image (montage). Since mosaic arrays are blended in such a way that the most recent data take precedence, a 24-hr. discontinuity occurs at the western edge of the most recent mosaic. Near the poles, where overlap is extreme, imagery is continually being replaced (see fig. 9). Both map montage arrays are maintained in auxiliary memory at all times and neither is purged of information. A continuous coverage product is thereby always available for image extraction or other use.

The polar map array is laid on the map base with the pole in the center and the 80° W. meridian aligned with the lower half of the center column. The Southern Hemisphere array is on a like map base except that the 80° W. meridian is the prime upper vertical meridian so

that the Southern Hemisphere array is merely an extension of the Northern. The Mercator montage, with a longitude extent of 410° and a latitude coverage from 35° N., to 27° S., has a grid square dimension of 13,120 columns by 2,000 grid rows. The 50° longitude extension precludes splitting an orbital mosaic. The starting position for blending can be selected at any longitude. There is also an image discontinuity at the lateral array boundary limits.

Both arrays are accessible in parts or segments so that one desiring a specific coverage is able to extract only data in the area of interest. Current products are being projected for operational use on a 6-hourly basis to meet synoptic deadlines (see figs.10 and 11). Variables which determine map coverage and resolution for the Mercator montage are set by alterable parameters. The polar montage array is bound to the NWP grid square system for compatibility purposes. A total of 830,000 (6-bit byte) computer storage bins is required for the larger segments of this rectification program and about 25 min. are required for processing one orbital picture swath.



Figure 8 - The full resolution Mercator mosaic uses part of the picture sequence shown in figure 6.

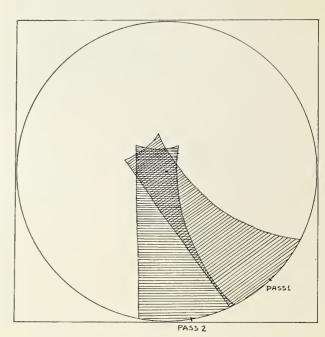


Figure 9 - Polar montage. Extreme overlap from pass to pass is indicated by the double shaded area.

CALIBRATION EFFECTS

In the description of the rectification process the mechanics of applying brightness calibrations were outlined. The data are altered according to the position in the picture-taking sequence of the frame in which the data reside and the position of the particular sample in the frame itself. The sample position is known only as it is a member of one of the 16X16 sample groups. This knowledge of how calibration is applied is now accompanied by an explanation of the reasons for calibrating and how the calibration function is obtained.

Apart from the cloud-earth scene, vidicon response variations may involve: (a) light transmission differences across the image plane arising from the properties of the optical system of the lens, (b) non-linearities in the x,y raster response pattern arising from the vidicon camera tube and associated electronics, (c) illumination variations due to solar angle (including extreme local sun glint problems), (d) "erase" efficiency and possible contamination



Figure 10 - Picture mosaics from adjacent passes are shown in polar stereographic montages. (Upper is Northern Hemisphere, lower is Souther Hemisphere) Reproduction uses a cathode ray film recorder. Latitude-longitude marks have been omitted from this early sample but the vertical seam between film segments is approximately 80° W.

from preceding scenes, and (e, time dependent anomalous responses - either response variations during a picture sequence from a single orbital pass or longer-period trends.

Based upon pre-launch calibration pictures and other information, item (d) is accepted as an uncontrolled contaminant with a brightness variation contribution of about 5 percent. If the camera system can be regarded as a stable sensor package, then one may attempt to compensate for the other listed items through a brightness calibration procedure.

The most glaring influence of the instrumentation on the data comes from the optical system (filter, aperture, and lens) immediately in front of the vidicon. The aperture controls the amount of light energy impinging on the vidicon face. The filter virtually eliminates the blue light, hence, the light scattered by the atmosphere. These two components are most relevent in determining absolute brightness response. The lens system employed introduces a substantial vignette effect as shown in figure 12. This effect produces a target-shaped, halftone pattern because the digitized data are discrete. Each digital level constitutes a zone. We describe each zone by the fraction (p;) of the true image brightness presented to the camera system. The inverse of this fraction is the calibration factor for the zone. Two or more camera systems can be "equalized" by establishing

the gradient function and maximum p; for each. The target patterns for the two cameras of ESSA 1 are illustrated in figure 13. The satellite is sun synchronous, assuring constant sun angles to the cameras. Camera 1 is pointed away from the sun and has a larger aperture setting than camera 2 which is pointed toward the sun. Although the biased aperture settings tend to equalize the overall brightness of the images presented to the vidicons, the application of the "target factors" from figure 12 is required to maintain the integrity of the object illumination. The ability to apply such factors efficiently is significantly unique to computer processing. There is thus a facility for linearizing the dynamic gray scale range which is not fully realized in conventional photo processing.

Individual pictures so adjusted are merged into an orbital mosaic as a homogeneous image.



Figure 11 - Adjacent mosaics have been combined into a multi-pass Mercator montage similar to figure 9. Early versions of the brightness calibration tables have been used in both figures.

Merging mosaics into a montage poses the peculiar problem of accounting for the difference between illumination patterns. The light energy flux at or near the earth's surface is a function only of the angle of incidence of the light ray (disregarding atmospheric refraction, absorption, and scattering). This means that there is a flux discontinuity where mosaics join. The flux is equalized among the mosaics by employing the inverse of the solar incident angle function as a factor. This alteration of the observed images must be considered when the data are examined. There are two major ramifications of this illumination normalization: (a) the angle of incidence is unchanged, leaving (in fact enhancing) the original shadows, and (b) such physical measures as albedo must be considered "normalized" by rendering the flux constant throughout. The original scene is recoverable through knowledge of the illumination patterns and the method of merging the mosaics.

Current ESSA satellites, being sun synchronous, have constant illumination patterns which, when combined with the vignette patterns, provide frame calibration patterns that are dependent on latitute. By means of a standard set of zone tables according to the latitude of the center point of the picture, both calibrations are applied at the same time. The "latitudes" of the reference array tables change with the seasonal position of the sun with respect to the earth's latitude-longitude system.

Although the calibrating scheme has been discussed conceptually, the actual calibration tables are determined empirically. This technique automatically includes compensation for time dependent anomalous response. The structuring of the zone tables permits calibration for any empirically determined nonlinearities.

8. SUMMARY REMARKS

Satellite picture data processing by computer has been started in a practical real time operation. Flexibility has been the watchword in the logical design of the programing effort. The system is thereby relatively insensitive to superficial differences in image source data and may easily be modified as developments in output products dictate. Quantized outputs are

expected to encourage further objective utilization of the products.

The programs represent a substantial investment in machine language coding so as to take maximum advantage of computer speed and memory capacity. Even so, projected higher resolution (and possibly color) imaging devices threaten to increase data volumes by one or two orders of magnitude within the next several years. This would likely lead to different applications and different information extraction goals. All of this could saturate today's largest computer and might well indicate the need for the newly projected parallel network computer.

Before that time, experience should provide guidelines for the specification of practical volumes of data to be computer processed commensurate with the application goals.



Figure 12 - A digital reproduction of a picture of a uniformly illuminated screen taken by the camera (AVCS) of the ESSA 3 prototype is shown. Although the screen illumination was maximum white, the camera response at the corners was near zero, showing the severity of the vignette effect.

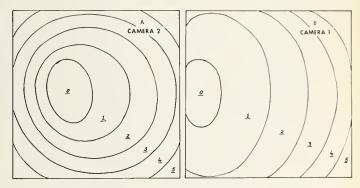


Figure 13 - The calibration zones are shown graphically for the two ESSA 1 cameras.

(A) Camera 2 points toward the sun and (B) camera 1 points away from the sun. The vignette effect is obvious. The shift of the centers of the patterns (zone zero) to the left demonstrates the influence of the sun's position which is to the left with respect to the figures. Zone zero is the area of maximum response, hence, least adjustment. As the zone number (index) increases the amount of calibration adjustment increases.

1218

12

43E

765 665 535 40S 27S 13S

ON 13N 26N 39N 52N 65N

119 092636

638

1368

1 1

1443

1.1

F

73E

125

084557

755 645 515 385 245 115

15N 28N 42N 54N 67N

เกร

218 040445

218 055931

1428

114E

1742

1743

8 2 W

9

2 V

765 665 535 405 275 145

765 655 535 405 275 135

ON 12N 25N

ON 12N 25N

38N 51N 64N

39N 52N 64N

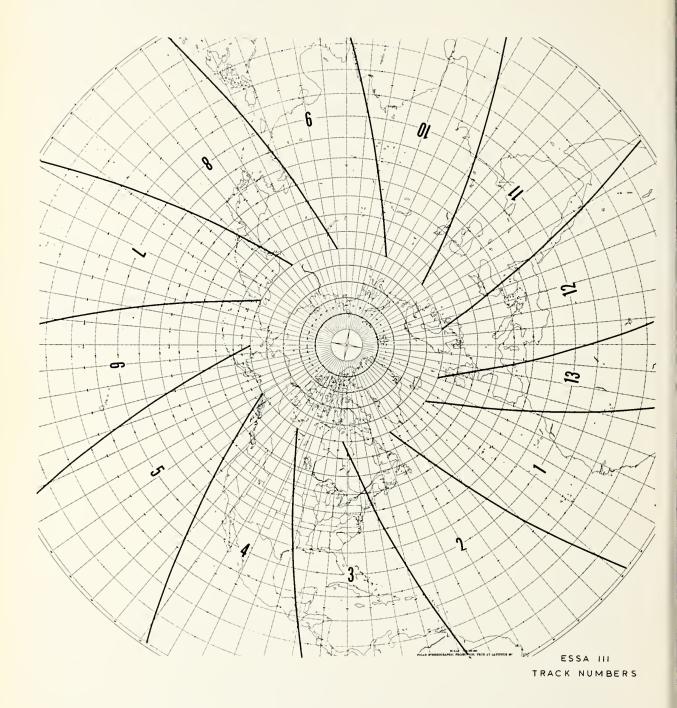
PASS	TRACK	С	s	ALON	HODA HRHINSE	SUB-SATELLITE POINT PICTURE LOCATION LAT LAT LAT LAT LAT LAT LAT LAT LAT
1969	9	1	W	1150	3 8 055054	745 635 505 375 245 115 2N 15N 28N 41N 54N 67N
1970	10	1	Ä	338	3 8 074544	735 625 495 365 235 105 2N 16N 29N 42N 55N 67N
1971	11	1	F	58C	3 8 094037 3 8 113533	735 615 495 365 235 95 3N 16N 30N 43N 56N 68N 725 615 485 355 225 85 4N 17N 30N 44N 57N 69N
1973	1	i	ř	OV	3 8 133013	725 615 485 355 225 95 4N 17N 30N 43N 56N 68N
1974	2	1	F	28V	3 8 152438	735 615 495 365 235 95 3N 16N 29N 42N 55N 68N
1975	3	1	Ē	57V	3 8 171905	735 625 495 365 235 105 2N 16N 29N 42N 55N 67N
1976	5	1	F	86V	3 8 191335	735 625 505 375 245 105 2N 15N 28N 42N 54N 67N 745 625 505 375 245 115 2N 15N 28N 41N 54N 67N
1978	6	i	F	143V	3 8 230245	745 635 505 375 245 115 2N 15N 29N 41N 54N 67N
1979	7	1	F	172W	3 9 005713	745 635 515 385 245 115 1N 14N 28N 41N 54N 66N
1980	8	1	V	1600	3 9 025153 3 9 044637	745 635 505 375 245 115 1N 14N 28N 41N 54N 66N 745 635 505 375 245 115 1N 15N 28N 41N 54N 67N
1982	10	1	V	131E 102E	3 9 064123	745 625 505 375 245 115 1N 15N 28N 41N 54N 67N 745 625 505 375 245 105 2N 15N 28N 42N 54N 67N
1983	11	i	v	740	3 9 083613	735 625 495 365 235 105 2N 16N 29N 42N 55N 67N
1984	12	1	F	458	3 9 103106	735 615 495 365 225 95 3N 16N 30N 43N 56N 68N
1985	13	1	ŗ	160	3 9 122544 3 9 142008	735 615 495 365 235 95 3N 16N 29N 42N 55N 68N 735 625 495 365 235 105 2N 15N 28N 42N 55N 67N
1987	1 2	1	F	12V	3 9 161435	73\$ 62\$ 49\$ 36\$ 23\$ 10\$ 2N 15N 28N 42N 55N 67N 74\$ 63\$ 50\$ 37\$ 24\$ 11\$ 1N 15N 28N 41N 54N 66N
1968	3	i	F	70V	3 9 180906	745 635 505 375 245 115 1N 14N 28N 41N 54N 66N
1989	4	1	ŗ	98\/	9 200339	745 635 515 385 255 115 1N 14N 27N 40N 53N 66N
1990	5	1	F	127V	3 9 215816	745 635 515 385 255 115 1N 14N 27N 40N 53N 66N 755 645 515 385 255 125 0N 14N 27N 40N 53N 65N
1992	6	1	v	156W	3 9 23524 5 310 014725	755 645 515 385 255 125 ON 14N 27N 40N 53N 65N 755 645 515 385 255 125 ON 14N 27N 40N 53N 66N
1993	8	i	ũ	147E	310 034209	745 635 515 385 255 125 1N 14N 27N 40N 53N 66N
1994	9	1	W	118E	310 053656	745 635 505 385 245 115 1N 14N 28N 41N 54N 66N
1995	10	1	F	90£	310 073146 310 092638	74\$ 63\$ 50\$ 37\$ 24\$ 11\$ 2N 15N 28N 41N 54N 67N 73\$ 62\$ 49\$ 36\$ 23\$ 10\$ 2N 16N 29N 42N 55N 68N
1997	12	i	F	320	310 032838	735 615 485 355 225 95 3N 17N 30N 43N 56N 68N
1998	13	i	F	40	310 131615	735 615 485 355 225 95 3N 16N 29N 43N 55N 68N
1999	1	1	F	25V	310 151039	735 625 495 365 235 105 2N 16N 29N 42N 55N 67N
2000	2	1	F	25V 54V	310 1 70506 310 1 85936	745 635 515 355 245 125 2N 15N 28N 42N 54N 67N 755 635 515 385 255 125 2N 14N 28N 41N 54N 67N
2002	4	1	F	1110	310 205410	745 635 505 375 245 115 1N 14N 28N 41N 54N 66N
2003	5	i	F	iiiw	310 224846	755 645 525 385 245 115 1N 14N 28N 41N 54N 66N
2004	6	1	F	168V		745 635 515 385 255 115 1N 14N 27N 41N 54N 66N
2005	7	1	V	163E	311 023805 311 043249	745 635 505 385 245 115 1N 14N 28N 41N 54N 66N 74S 635 505 375 245 115 1N 15N 28N 41N 54N 67N
2007	9	1	V	1068	311 062736	745 625 505 375 245 115 14 154 264 414 544 674
2008	10	1	V	776	311 082226	735 625 495 365 235 105 2N 16N 29N 42N 55N 68N
2009	11	1	F	48E	311 101718	735 615 495 365 225 95 3N 16N 30N 43N 56N 68N
2010	12	1	F	20E	311 121157 311 1 4062 1	735 615 495 365 235 95 3N 16N 29N 42N 55N 68N 735 625 495 365 235 105 2N 15N 29N 42N 55N 67N
2012	2	i	F	38√		745 635 505 375 245 115 1N 15N 28N 41N 54N 66N
2013	3	1	F	66V		745 635 505 375 245 115 1N 14N 28N 41N 54N 66N
2014	4	1	ŗ	95V		745 635 515 385 255 115 1N 14N 27N 40N 53N 66N
2015	5	1	F	124W 152W		745 635 515 385 255 125 1N 14N 27N 40N 53N 66N 75S 645 515 385 255 125 0N 14N 27N 40N 53N 66N
2017	7	i	·	179€	312 013342	745 635 515 385 255 125 1N 14N 27N 40N 53N 66N
2018	8	1	¥	150€	312 032825	745 635 515 385 255 115 1N 14N 27N 41N 54N 66N
2019	9	1	V	1220	312 052312	745 635 505 375 245 115 1N 15N 28N 41N 54N 66N
2020	10	1	F	93E 64E	312 071802 312 091255	745 625 505 375 245 105 2N 15N 28N 42N 55N 67N 735 625 495 365 235 105 3N 16N 29N 42N 55N 68N
2022	12	i	F	360	312 031233	725 615 485 355 225 95 4N 17N 30N 43N 56N 69N
2023	13	1	F	75	312 130229	725 615 485 355 225 95 3N 16N 3ON 43N 56N 68N
2024	1	1	F	22V		735 625 495 365 235 105 2N 16N 29N 42N 55N 67N 735 625 505 375 245 105 2N 15N 28N 42N 54N 67N
2025	2	1	F	50V 79V		745 635 505 375 245 105 2N 15N 2BN 42N 54N 67N
2027	4	i	F	108W		745 635 505 375 245 115 1N 14N 28N 41N 54N 66N
2028	5	1	F	136V		745 635 505 375 245 115 1N 14N 28N 41N 54N 66N
2029	6	1	<u>.</u>	165W		745 635 515 385 255 115 1N 14N 27N 41N 54N 66N 74S 635 505 385 245 115 1N 14N 28N 41N 54N 66N
2030	7 8	1	Z	166E	313 022417 313 041900	745 635 505 385 245 115 1N 14N 28N 41N 54N 66N 74S 635 505 375 24S 11S 1N 15N 28N 41N 54N 67N
2032	9	i	ū	1098	313 061347	745 625 505 375 245 115 2N 15N 2BN 42N 55N 67N
2033	10	1	W	906	313 080836	735 625 495 365 235 105 2N 16N 29N 42N 55N 67N
2034	11	1	F	52E	313 10032 9 313 115807	735 615 495 365 225 95 3N 16N 30N 43N 56N 68N 735 615 495 365 235 105 3N 16N 29N 42N 55N 68N
2035	12	1	F	23E 6V		735 625 495 375 235 105 2N 15N 28N 42N 55N 67N
2037	2	i	F	34W		745 635 505 375 245 115 1N 15N 28N 41N 54N 66N
2038	3	1	F	63W		745 635 505 385 245 115 1N 14N 27N 41N 54N 66N
2039	4	1	F	92V		745 635 515 385 255 125 1N 14N 27N 40N 53N 66N 755 645 515 385 255 125 0N 13N 27N 40N 53N 65N
2040	5 6	1	F	120V		755 645 515 385 255 125 ON 13N 27N 40N 53N 65N 755 645 515 385 255 125 ON 13N 27N 40N 53N 65N
2042	7	i	F	178V		
2043	8	1	¥	154E	314 031430	

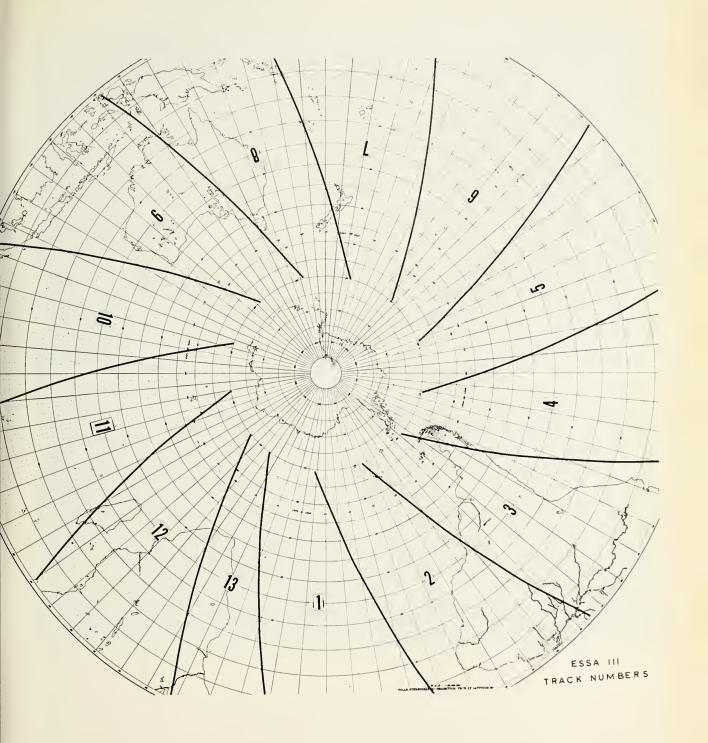
PASS T	HACK	C	5	ALON	MOUA	HKMNSL	NSE SUB-SATELLITE POINT PICTURE LOCATION											
							LAT	LAT	LAT	LAT	LAT	LAT	LAT	LAT	LAT	LAT	LAT	LAT
2269	8	2	V	155E	4 1	030502	755	655	525	395	265	135	114	14N	28N	41N	54N	66N
2770	9	2	V	1268	4 1	045948	755	645	515	385	255	125	1 %	15N	28N	41N	54N	67N
2271	10	2	F	976	4 1	065438	755	645	525	385	255	125	214	15N	28N	4214	55N	67N

TABLE 1
TRACK TABLE FOR ESSA III

Range of Ascending Nodes	Track No.
OW - 28.7W	1
28.7W - 57.4W	2
57.4W - 86.1W	3
86.1W - 114.8W	4
114.8W - 143.5W	5
143.5W - 172.2W	6
172.2W - 159.1E	7
159.1E - 130.4E	8
130.4E - 101.7E	9
101.7E - 73.0E	10
73.0E - 44.3E	11
44.3E - 15.6E	12
15.6E - 0	13

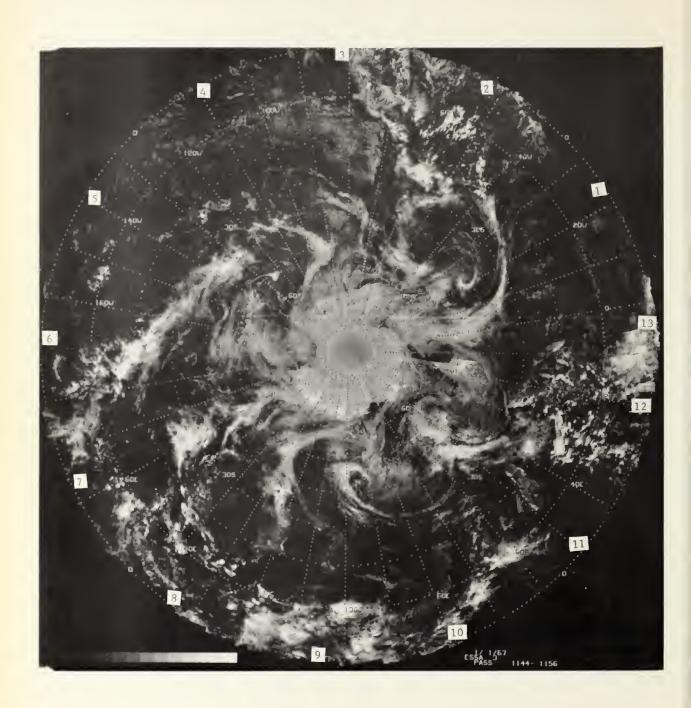
Difference between Tracks - 28.7°

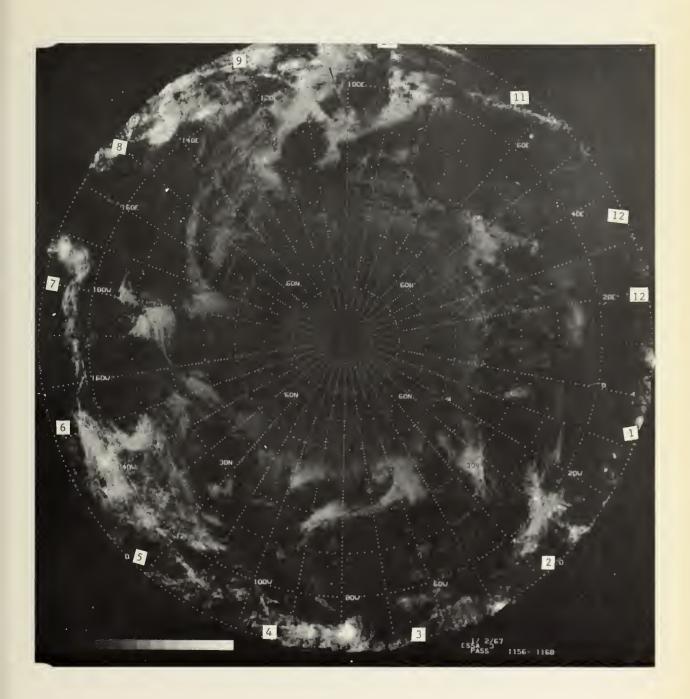


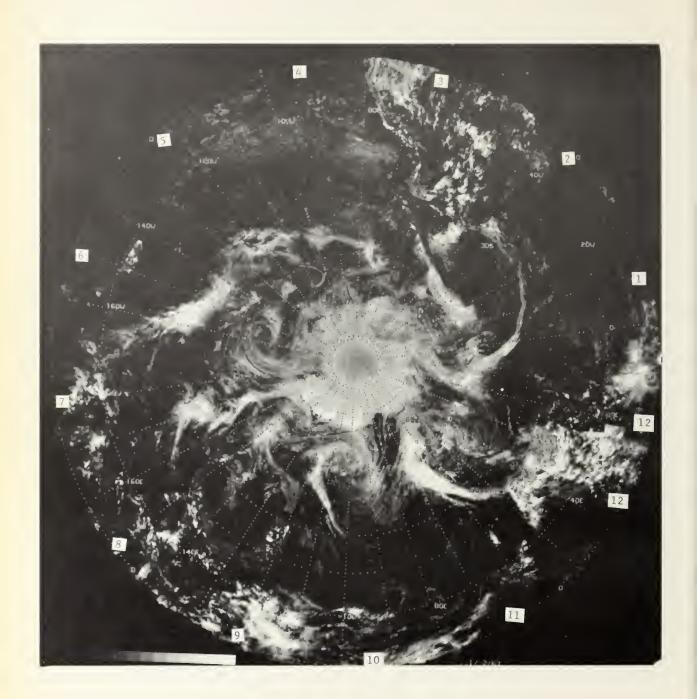




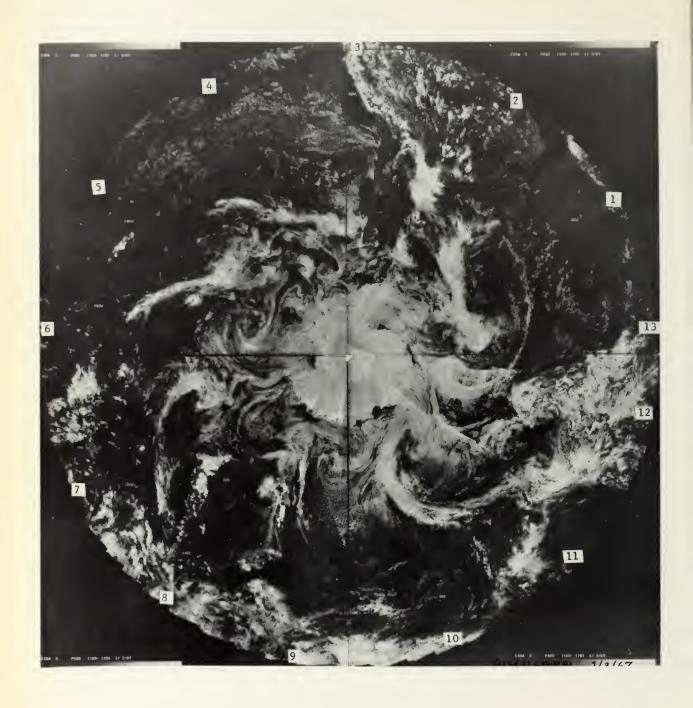
















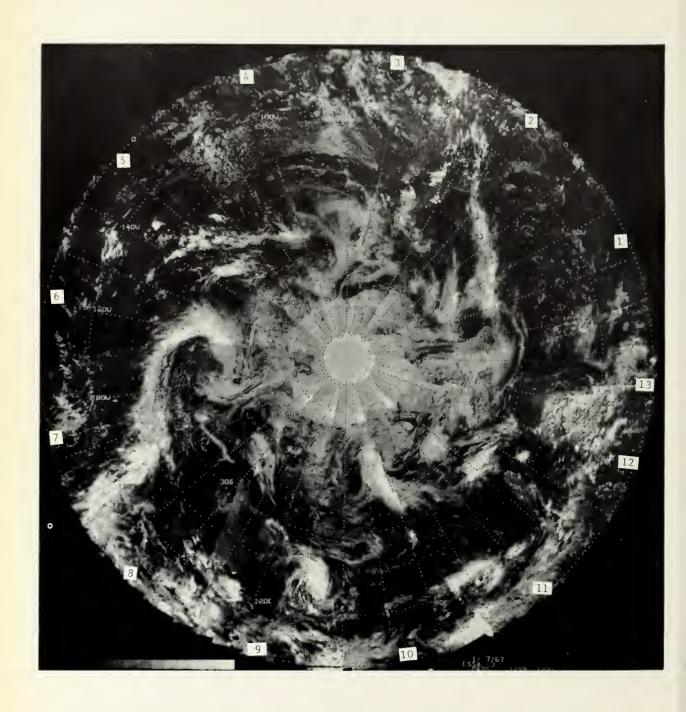




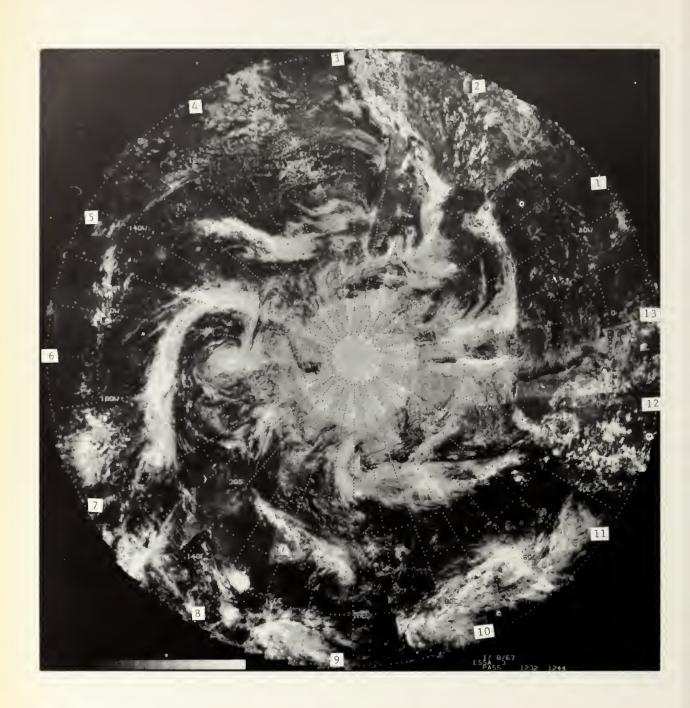












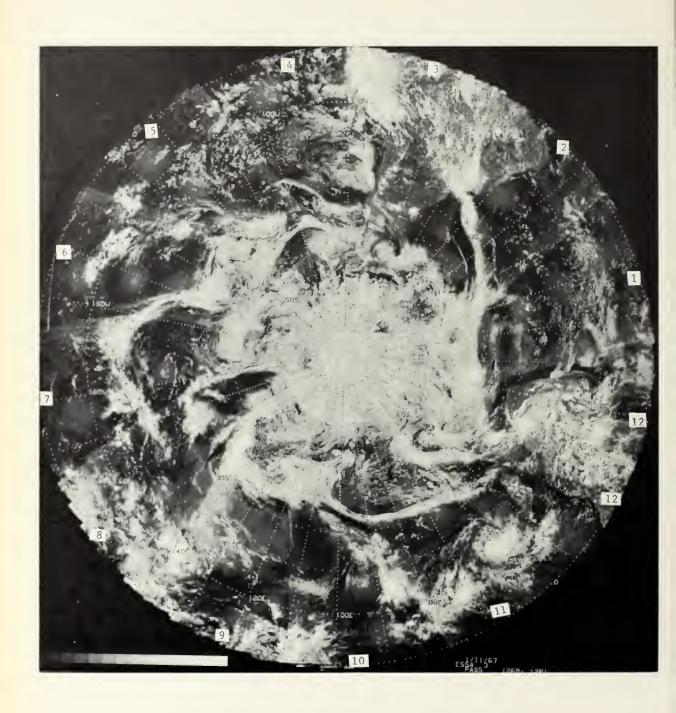


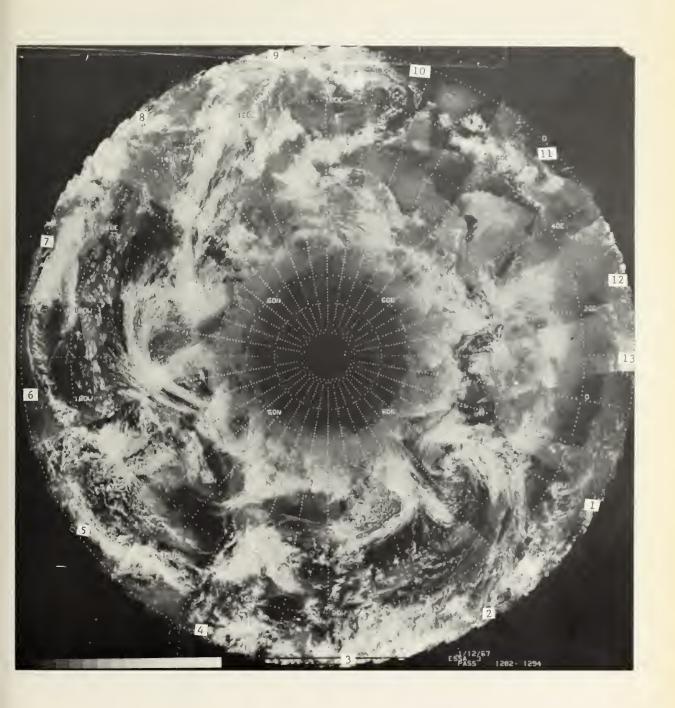


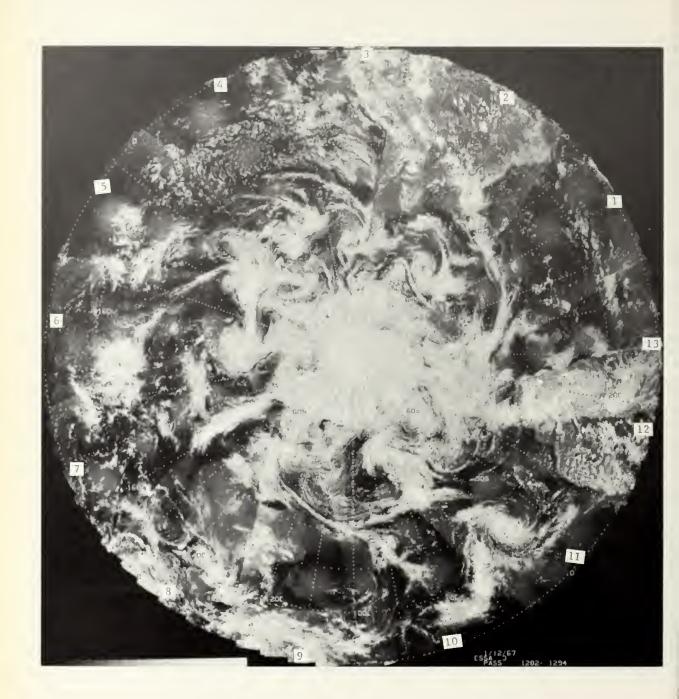




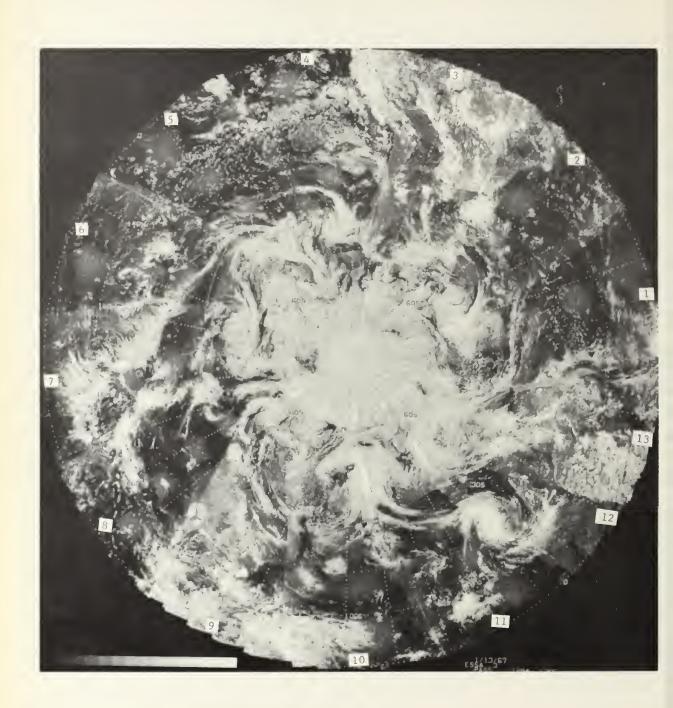




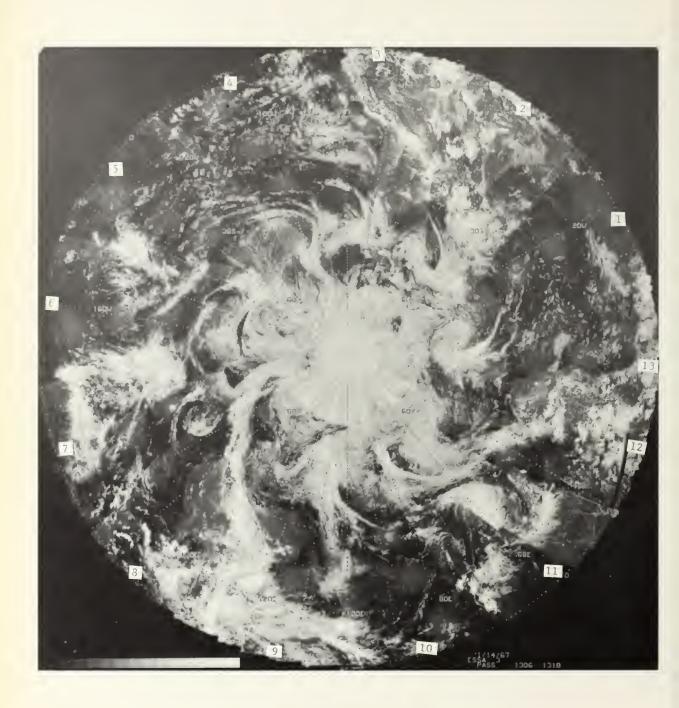






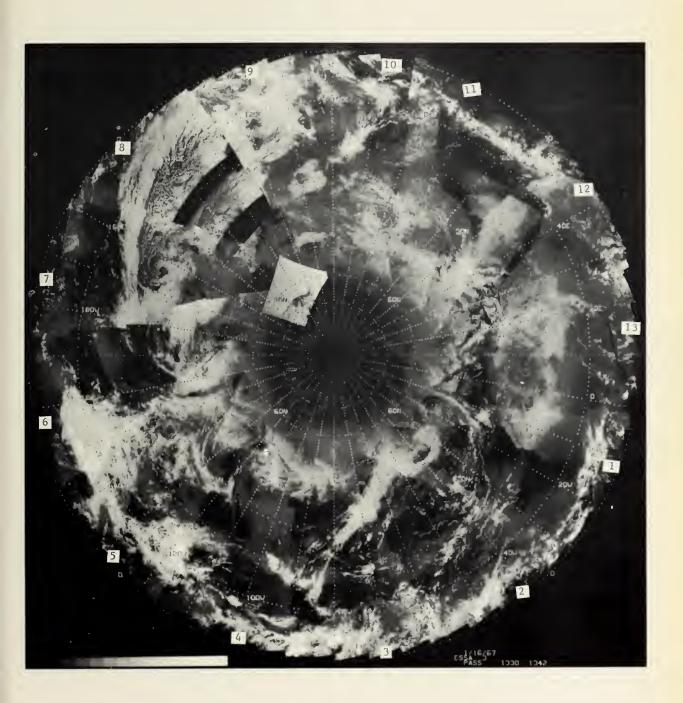














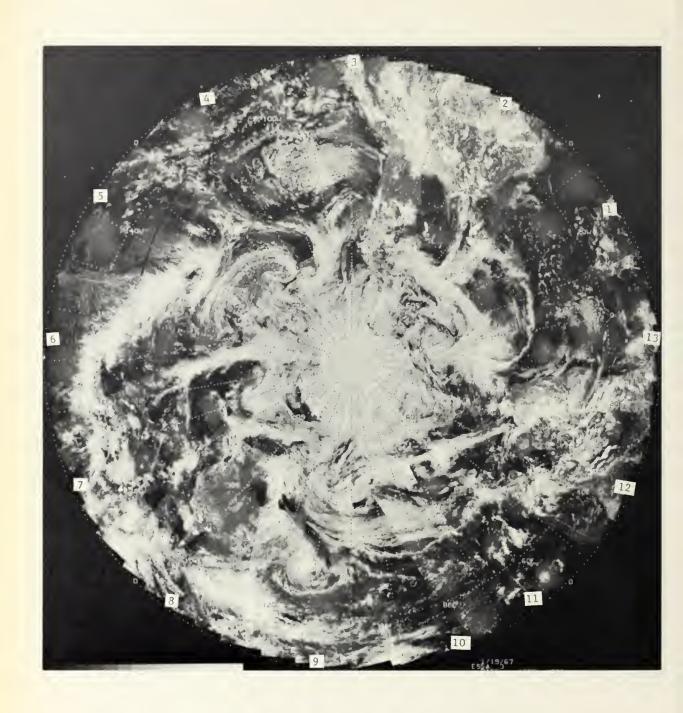




















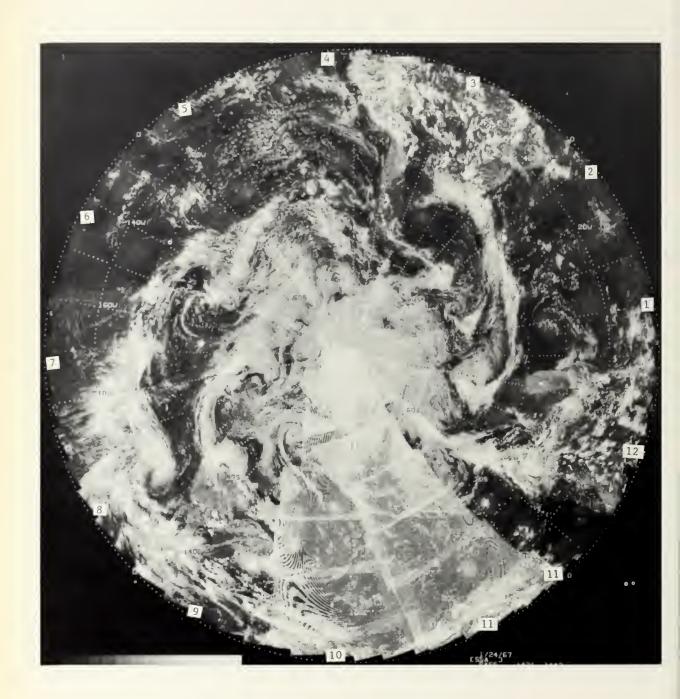


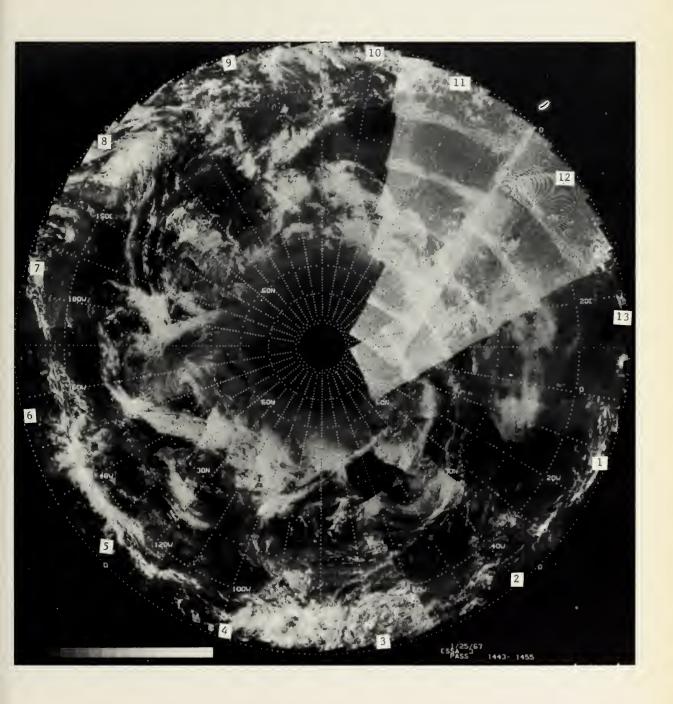


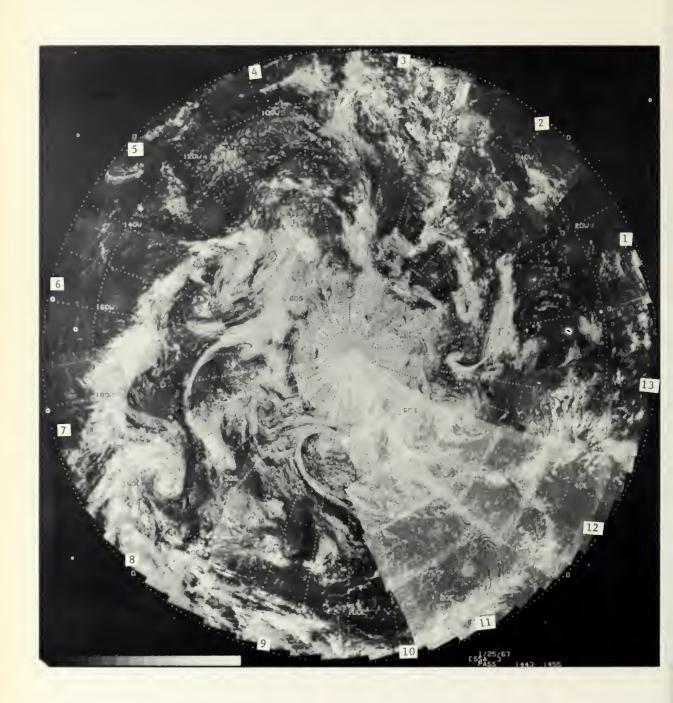


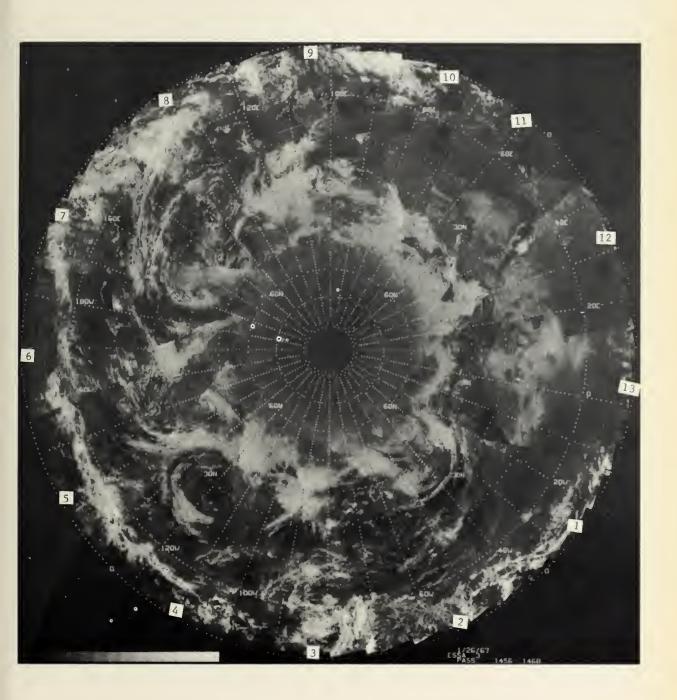


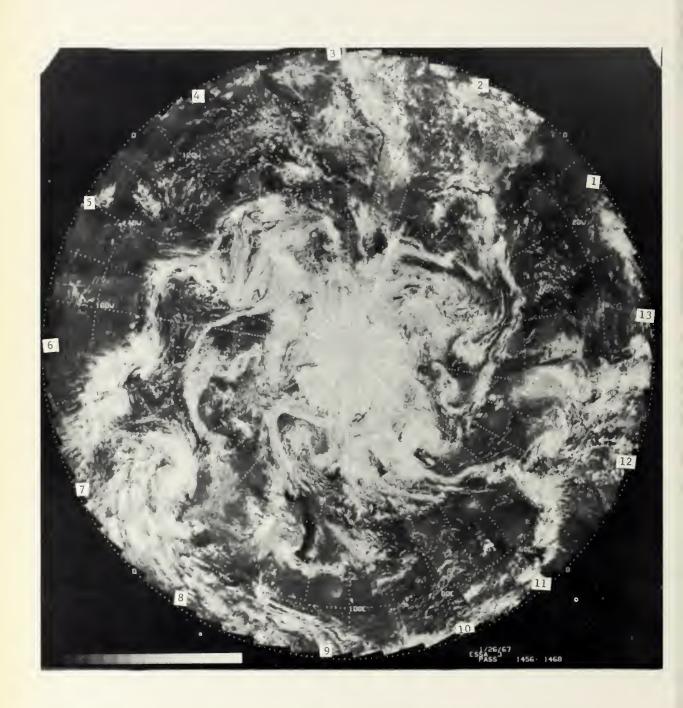


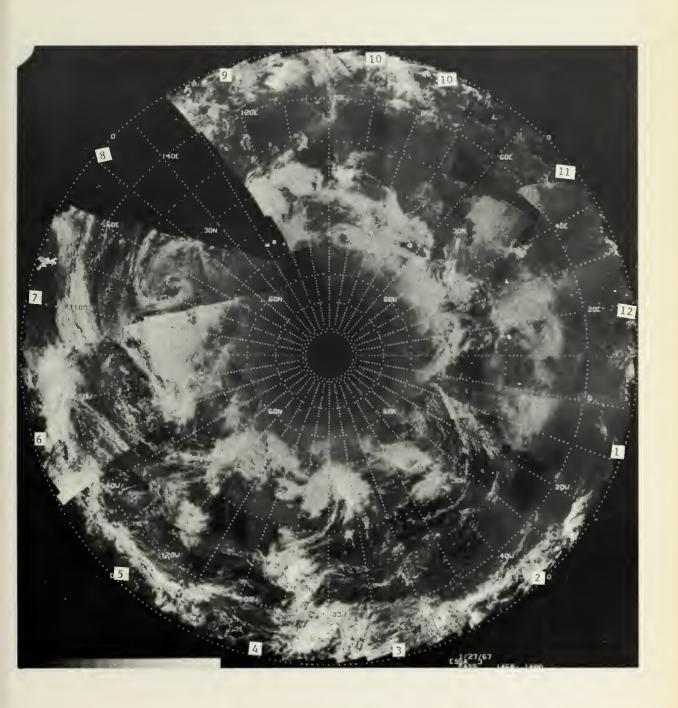




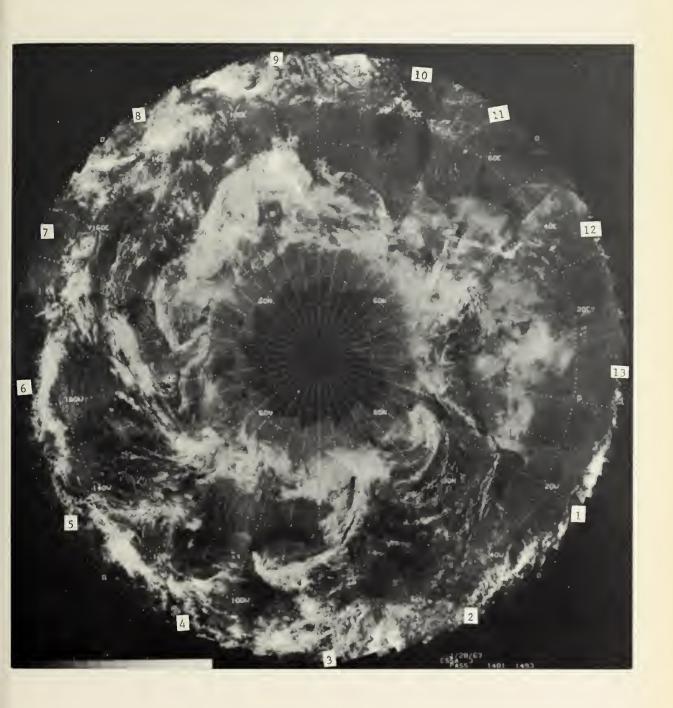


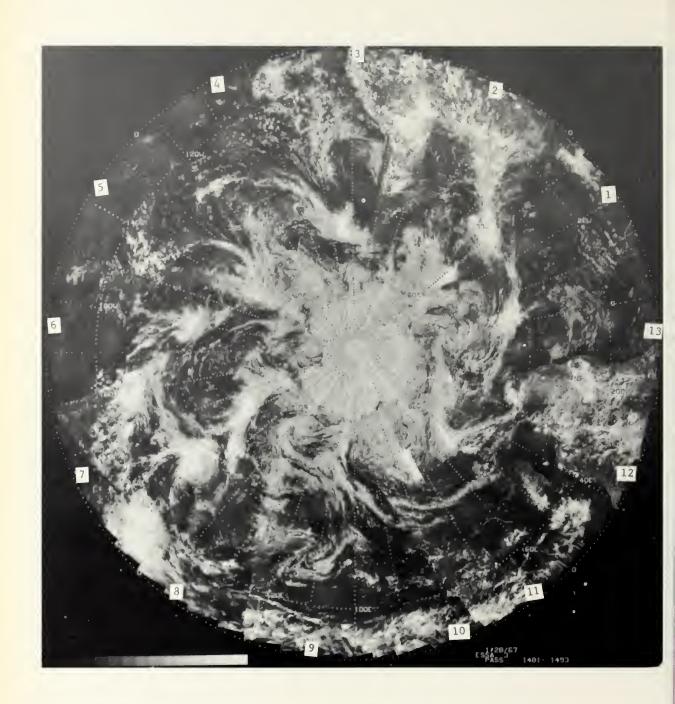




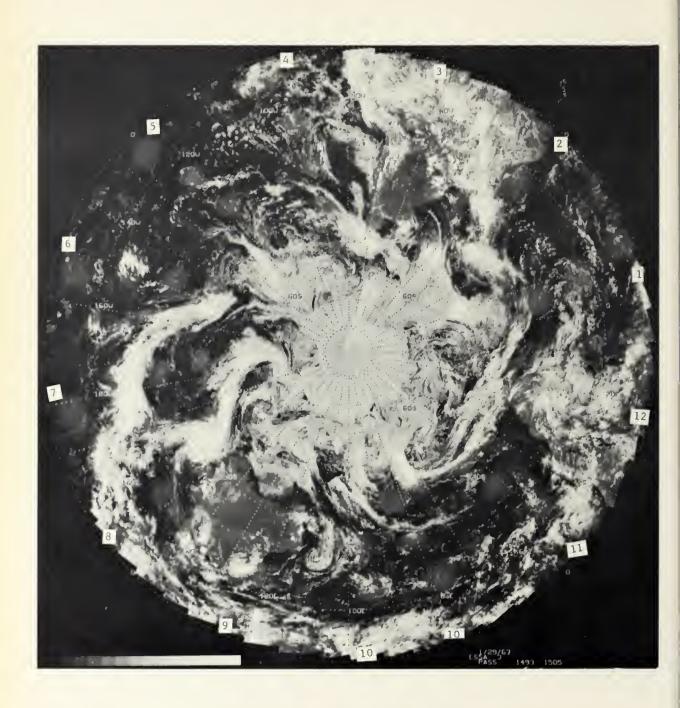


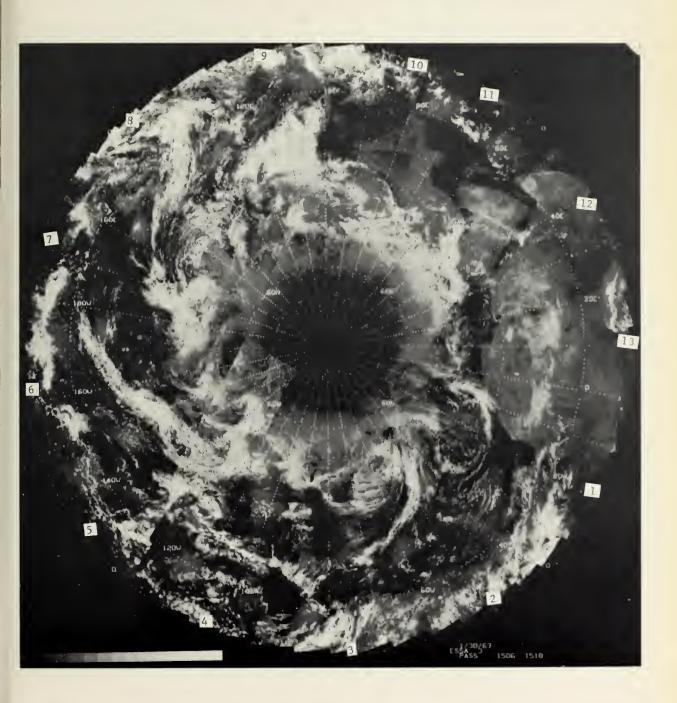


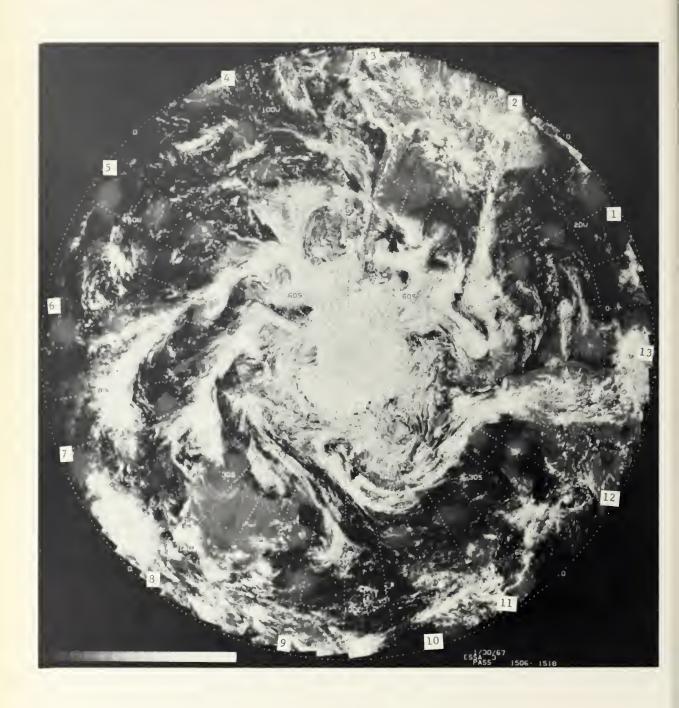


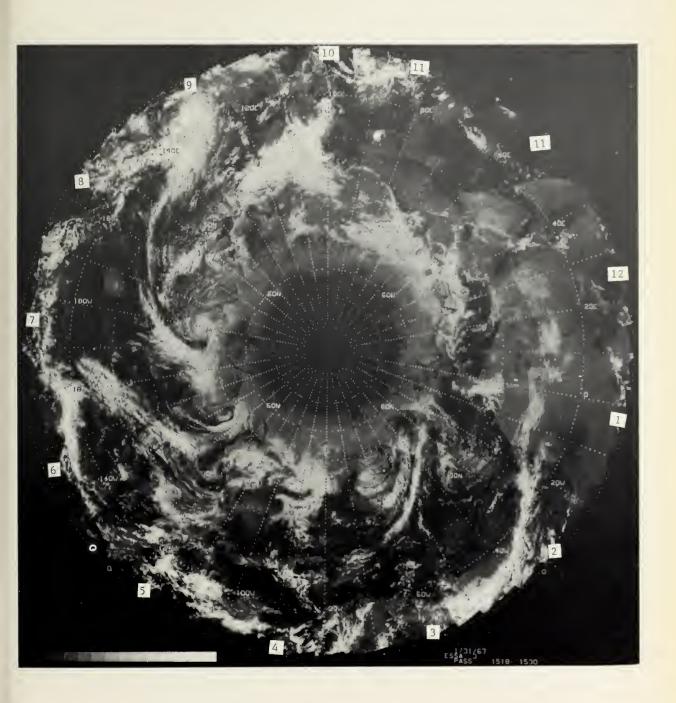




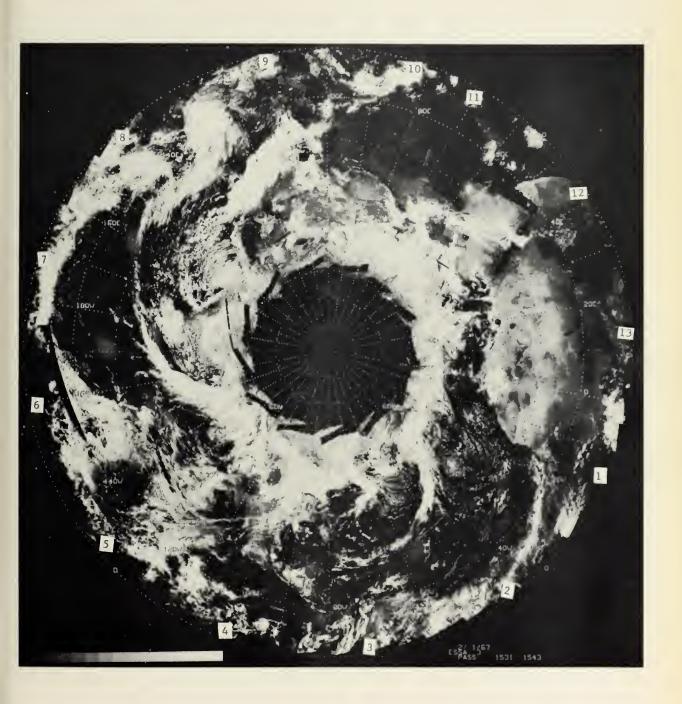


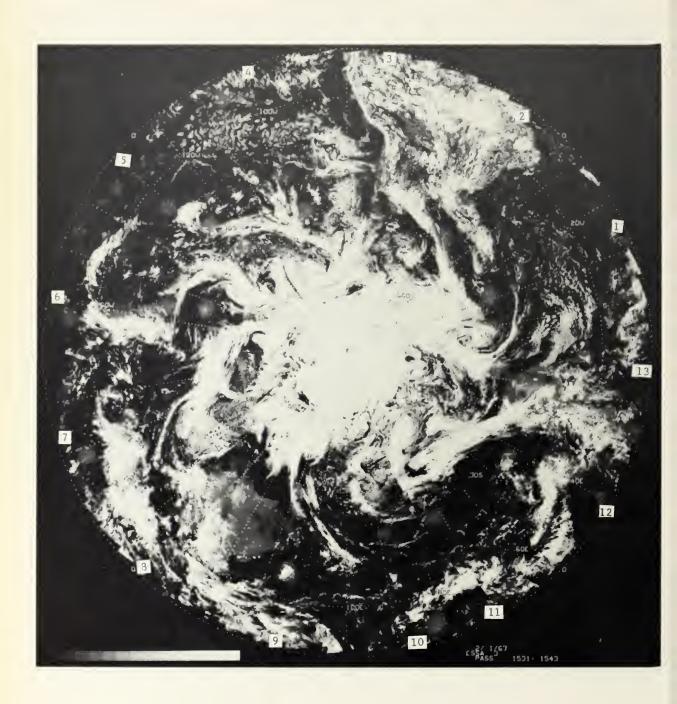






















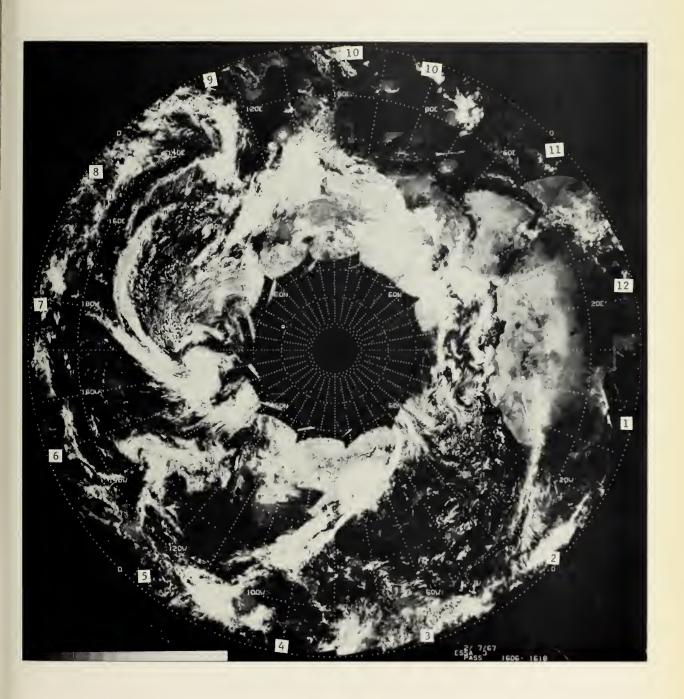








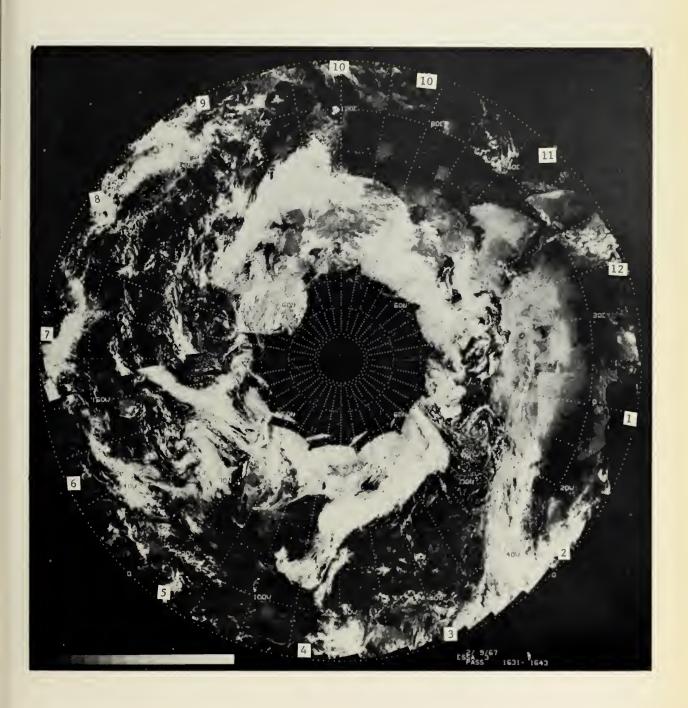




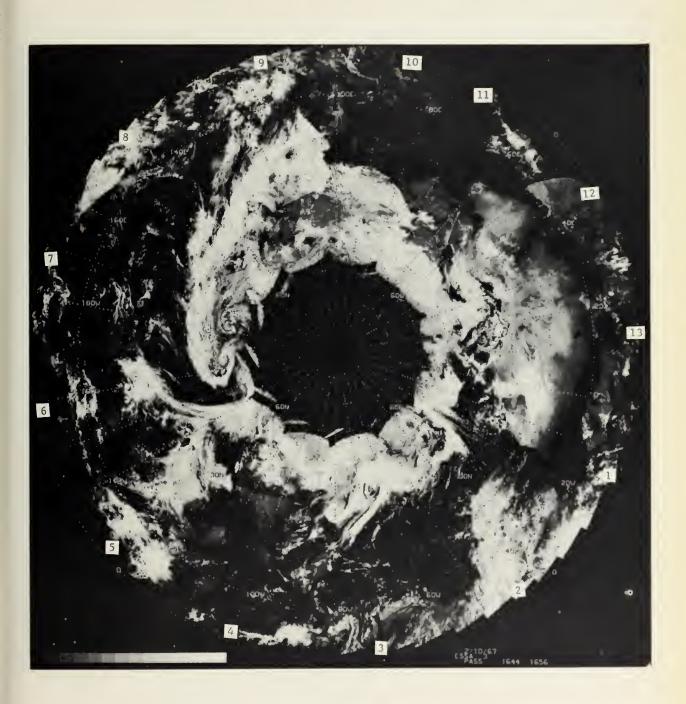




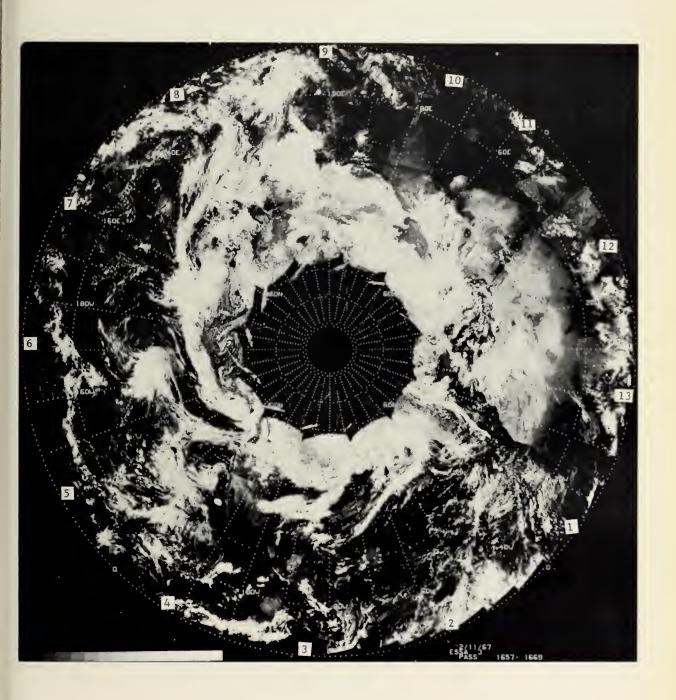
















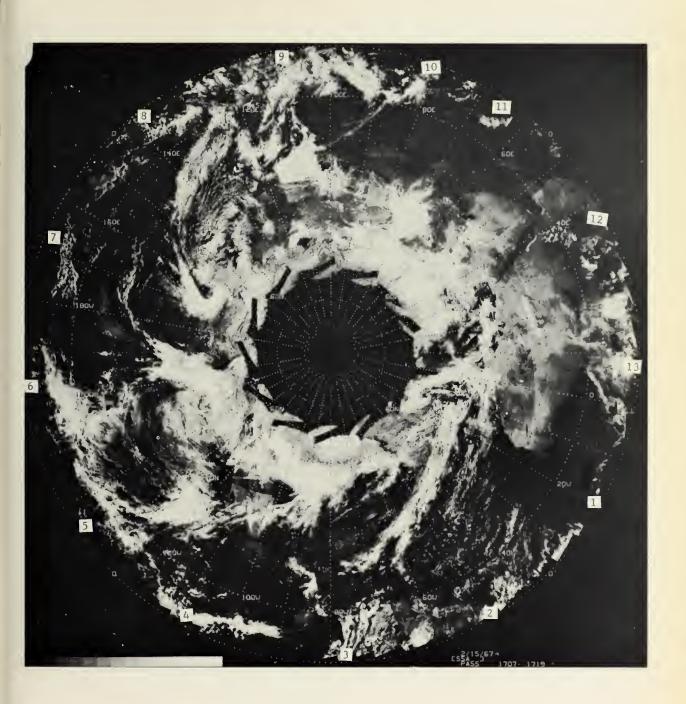




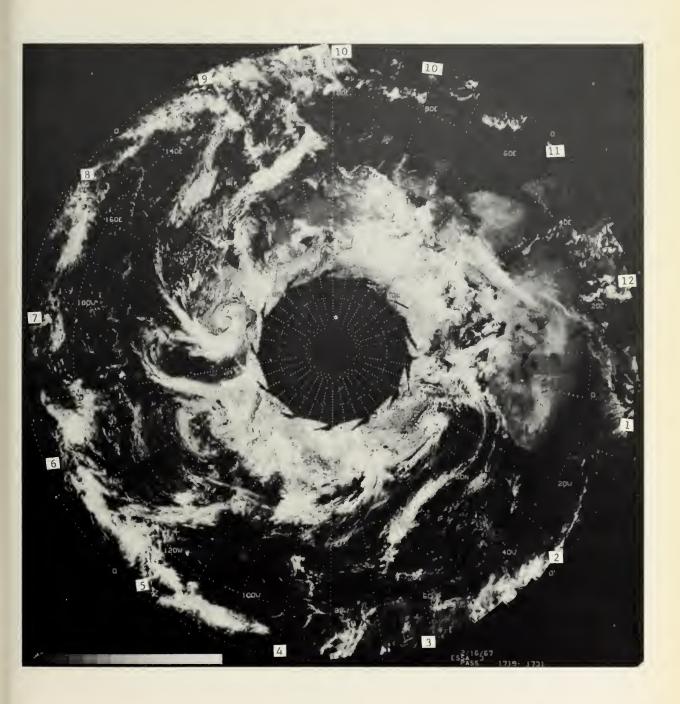




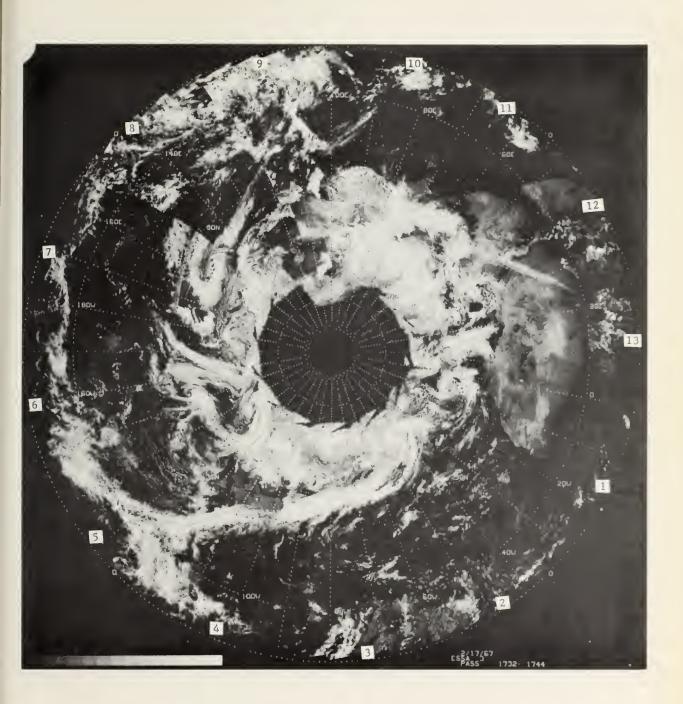


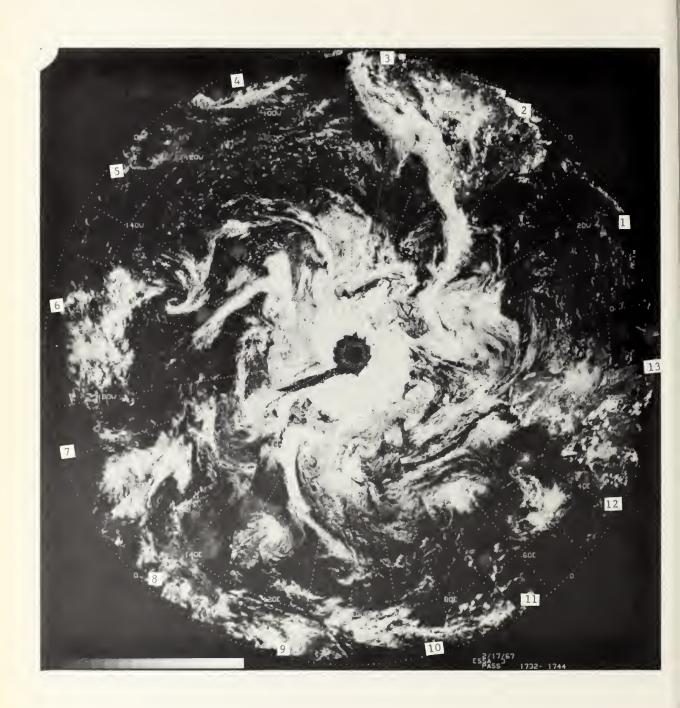


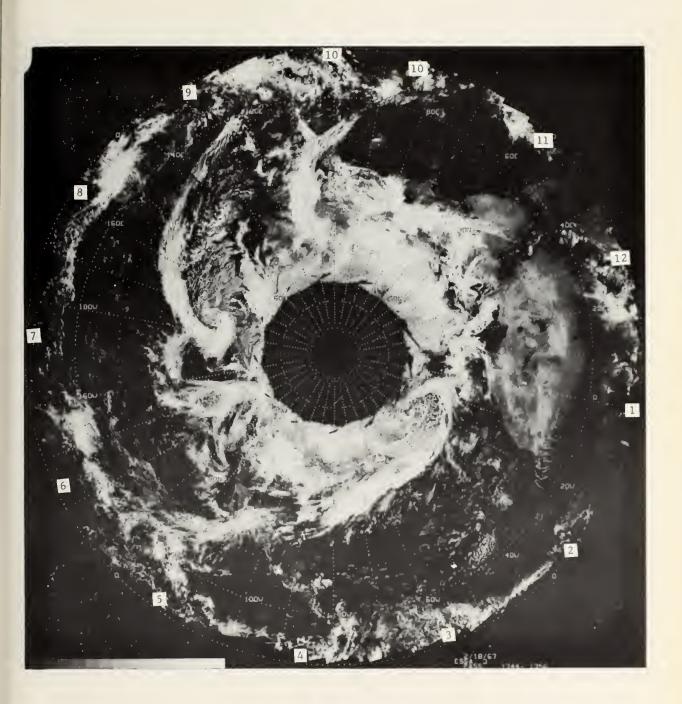








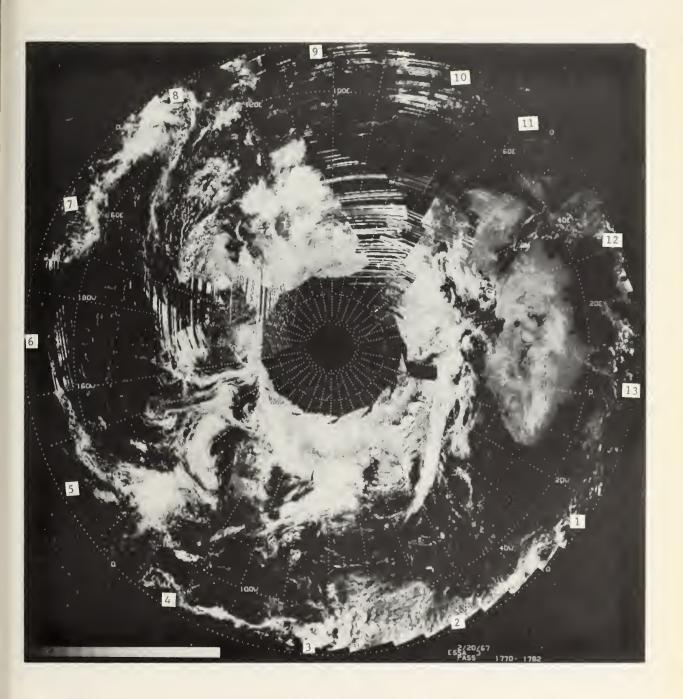








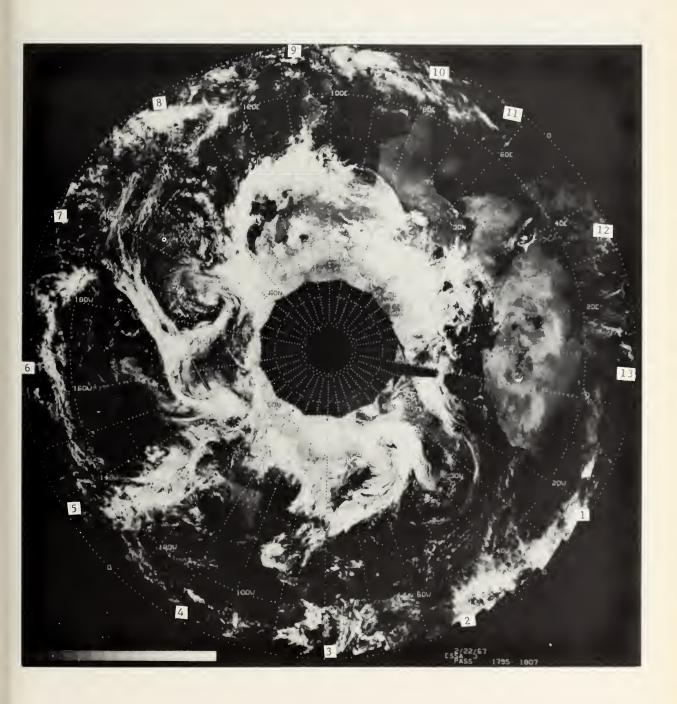
















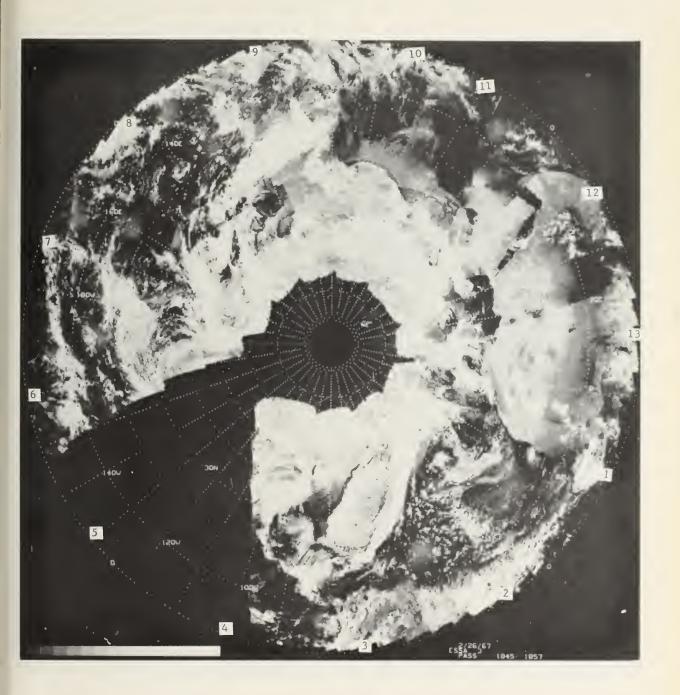








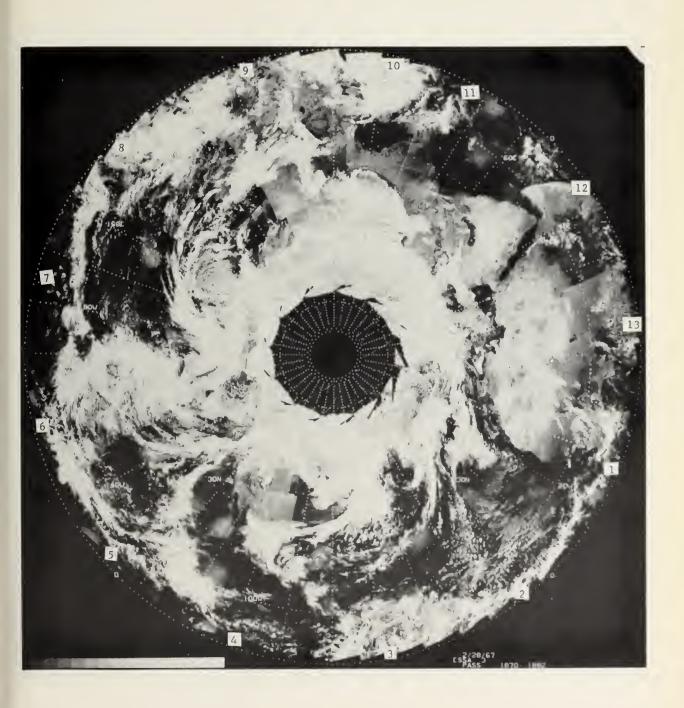
























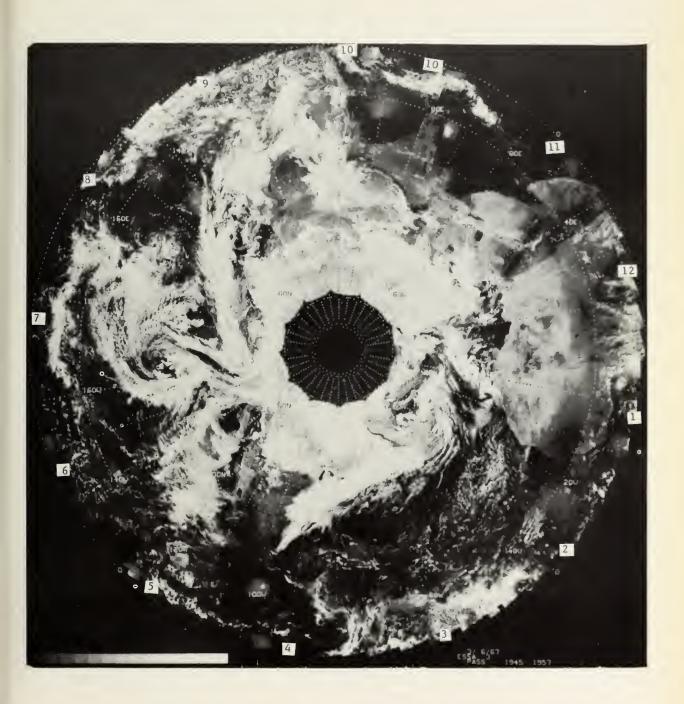
















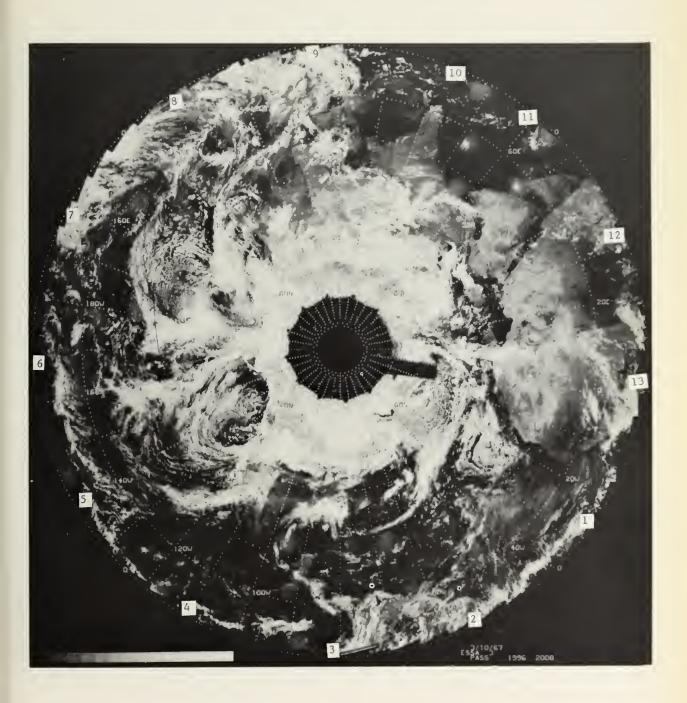




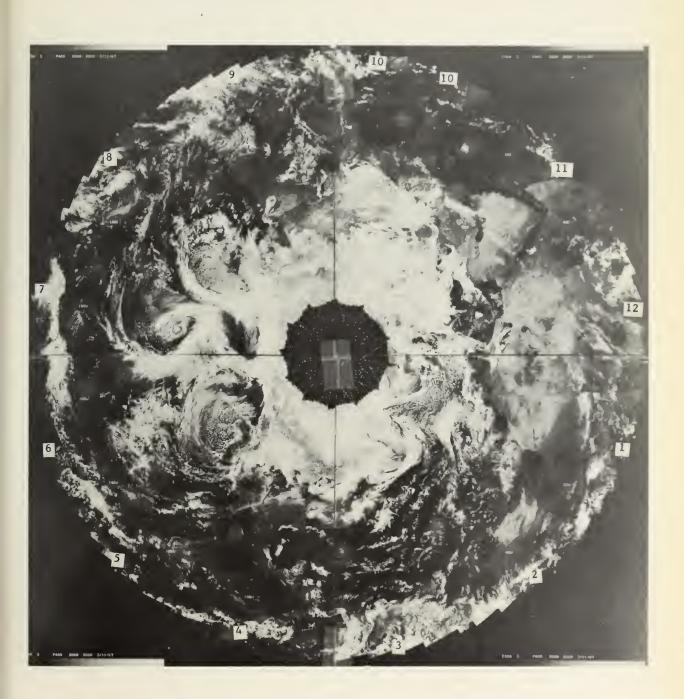




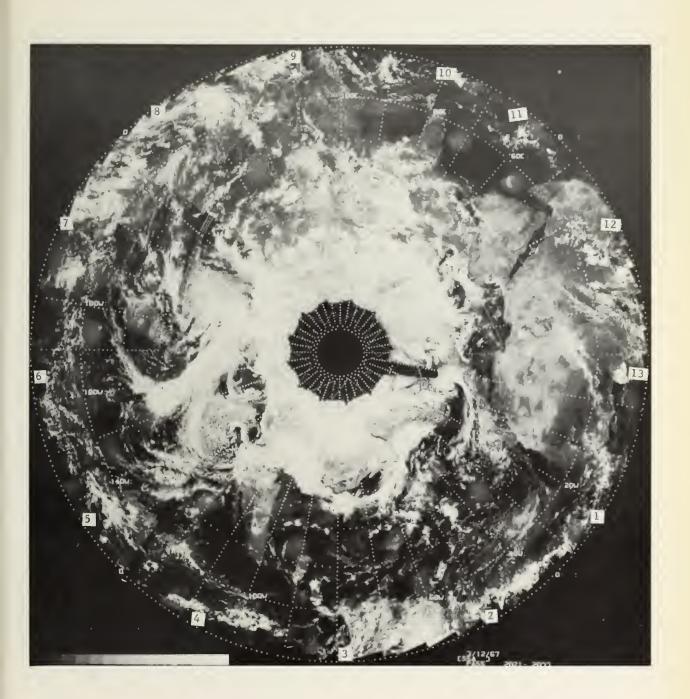


























































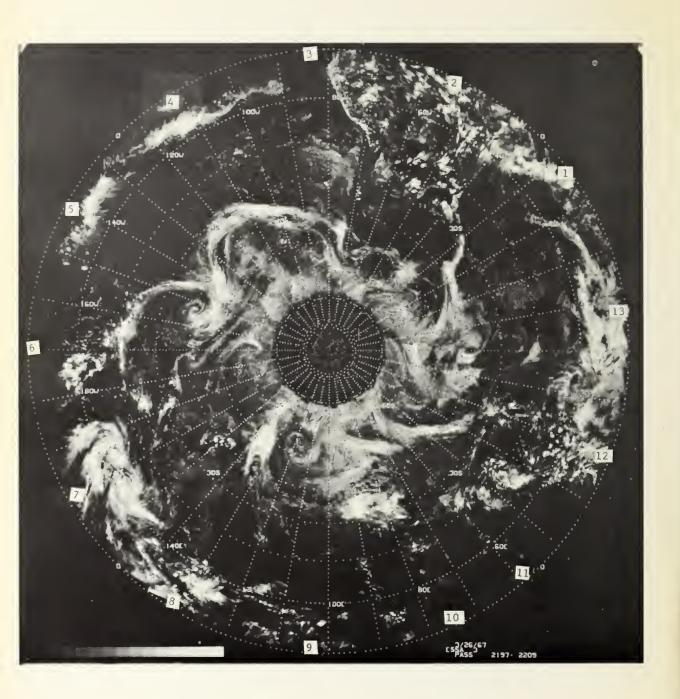


















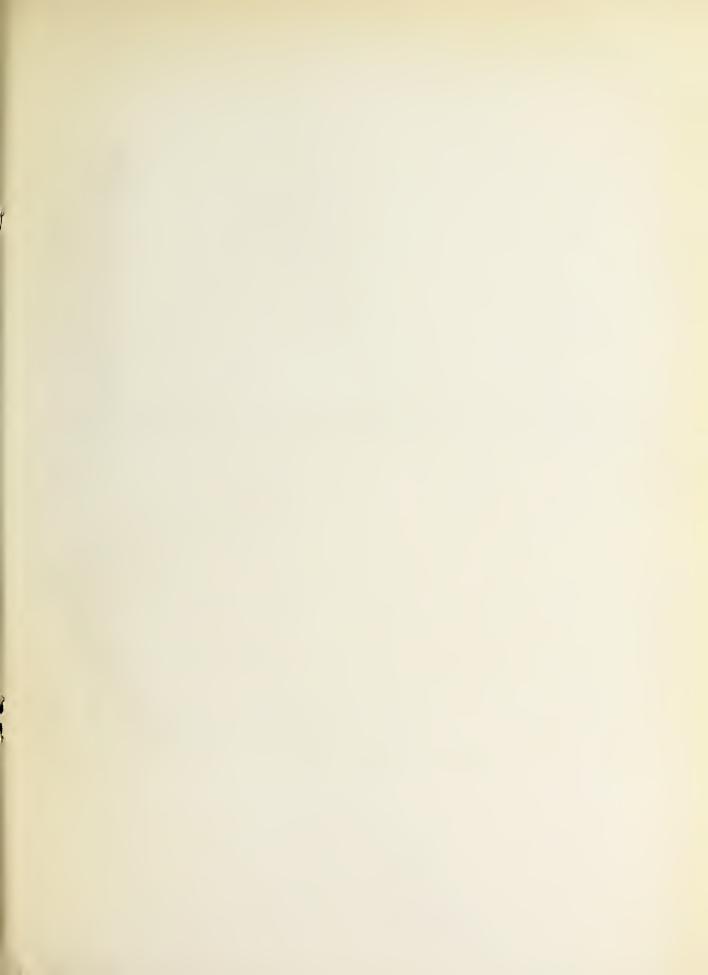












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